

Citation for published version:

Hamer, SL, Edge, AC, Swinbank, AM, Wilman, RJ, Combes, F, Salomé, P, Fabian, AC, Crawford, CS, Russell, HR, Hlavacek-Larrondo, J, McNamara, B & Bremer, MN 2016, 'Optical Emission Line Nebulae in Galaxy Cluster Cores 1: The Morphological, Kinematic and Spectral Properties of the Sample', *Monthly Notices of the Royal Astronomical Society*, vol. 460, no. 2, pp. 1758-1789. <https://doi.org/10.1093/mnras/stw1054>

DOI:

[10.1093/mnras/stw1054](https://doi.org/10.1093/mnras/stw1054)

Publication date:

2016

Document Version

Peer reviewed version

[Link to publication](#)

This is a pre-copy-edited, author-produced version of an article accepted for publication in *Monthly Notices of the Royal Astronomical Society* following peer review. The version of record Hamer, SL, Edge, AC, Swinbank, AM, Wilman, RJ, Combes, F, Salomé, P, Fabian, AC, Crawford, CS, Russell, HR, Hlavacek-Larrondo, J, McNamara, B & Bremer, MN 2016, 'Optical Emission Line Nebulae in Galaxy Cluster Cores 1: The Morphological, Kinematic and Spectral Properties of the Sample' *Monthly Notices of the Royal Astronomical Society*, vol. 460, no. 2, pp. 1758-1789 is available online at: <https://academic.oup.com/mnras/article/460/2/1758/2608996>

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Optical Emission Line Nebulae in Galaxy Cluster Cores 1: The Morphological, Kinematic and Spectral Properties of the Sample

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Accepted 8 March 2016. Received 26 February 2016; in original form 29 August 2014

ABSTRACT

We present an Integral Field Unit survey of 73 galaxy clusters and groups with the *Visible Multi Object Spectrograph* (VIMOS) on VLT. We exploit the data to determine the H α gas dynamics on kpc-scales to study the feedback processes occurring within the dense cluster cores. We determine the kinematic state of the ionised gas and show that the majority of systems ($\sim 2/3$) have relatively ordered velocity fields on kpc scales that are similar to the kinematics of rotating discs and are decoupled from the stellar kinematics of the Brightest Cluster Galaxy. The majority of the H α flux ($> 50\%$) is typically associated with these ordered kinematics and most systems show relatively simple morphologies suggesting they have not been disturbed by a recent merger or interaction. Approximately 20% of the sample (13/73) have disturbed morphologies which can typically be attributed to AGN activity disrupting the gas. Only one system shows any evidence of an interaction with another cluster member. A spectral analysis of the gas suggests that the ionisation of the gas within cluster cores is dominated by non stellar processes, possibly originating from the intracluster medium itself.

Key words:

galaxies: clusters: general - galaxies: elliptical and lenticular, cD - (galaxies:) cooling flows

1 INTRODUCTION

One of the key issues for our understanding of galaxy formation and evolution is the mechanism through which feedback from a galaxy affects the cooling of gas surrounding it. Simulations of galaxy formation which ignore non-gravitational

heating (e.g. Katz & White 1993) produce a galaxy population with an excess of massive galaxies when compared to the observed Universe. This occurs as a result of gas cooling being constrained only by its density, which results in rapid cooling and an over production of stars in the dense clumps where galaxies form. To address this problem the injection of energy into the gas from non-gravitational processes within the galaxies is invoked, a process dubbed ‘feedback’. By in-

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cluding feedback in the simulations the over production of massive galaxies is stopped resulting in a galaxy population in closer agreement to that which is observed (Bower et al. 2006; Croton et al. 2006).

While feedback is widely accepted as the means to slow the cooling of gas, the form that it takes is a highly debated topic. At face value X-ray observations of the central regions of massive galaxy clusters show intense X-ray emission suggesting that the Intra Cluster Medium (ICM) is undergoing significant radiative cooling (Fabian et al. 1981). This rapidly cooling gas should condense into cold gas clouds and/or form stars on short timescales relative to the age of the cluster (Fabian 1994). However, the cold gas mass (Braine & Dupraz 1994; McNamara & Jaffe 1994; O’Dea et al. 1994; Edge 2001; Salomé & Combes 2003) and star formation rates observed (McNamara & O’Connell 1989) are too low to be consistent with the mass of gas that should be cooling from the ICM.

Early observations with XMM-Newton’s highly sensitive Reflection Grating Spectrometer (RGS) failed to detect the X-ray spectral features of gas cooling at ~ 1 keV (Peterson et al. 2001; Tamura et al. 2001; Peterson et al. 2003). The apparent lack of gas at X-ray temperatures of < 1 keV suggests that some process is truncating the cooling to lower temperatures and preventing much of the gas from cooling further (see the review by Peterson & Fabian 2006). However, as the ICM core is still radiating away most of its energy through X-ray emission at a high rate there must be some process acting in the cores of cooling flow clusters which is continually reheating the gas. Thus “feedback” from non-gravitational processes needs to be invoked in order to inject energy into the ICM and balance the effects of X-ray cooling. Many possible contributors to ICM heating have been suggested such as starbursts (Veilleux et al. 2005), shocks from mergers (Markevitch & Vikhlinin 2007), sloshing of gas (ZuHone et al. 2010), conduction from the surrounding ICM (Voigt & Fabian 2004) and mechanical or radiative feedback from an AGN (McNamara et al. 2005; McNamara & Nulsen 2007).

For the most rapidly cooling cluster cores, the BCGs ubiquitously exhibit optical line emission (Heckman et al. 1989; Cavagnolo et al. 2008). Crawford et al. (1999) found significant line emission in 32 per cent of a sample of 201 BCGs selected from the Brightest Cluster Survey (BCS) X-ray selected sample (Ebeling et al. 1998). This line emitting gas at 10^4 K traces filamentary structures around the BCG (Hatch et al. 2005, 2006; McDonald et al. 2011) and direct comparison has shown qualitatively similar structures in the 10^7 K X-ray (Fabian et al. 2008) and 30 K molecular (Salomé et al. 2011) gas. This structural similarity suggests that the gas phases are linked with gas from the ICM cooling through the warm phase before quickly condensing out into cold molecular gas clouds. However, the masses of the gas at these temperatures are not consistent with that predicted from cooling (Johnstone et al. 1987). Once the gas is in molecular clouds, ionising radiation from within the cluster core can easily re-ionise and excite the surface of the gas clouds allowing the ionised gas to act as a proxy for the molecular gas.

Early studies of the brightest handful of BCGs showed a variety of complex spatial and velocity structures present in the ionised gas (Crawford et al. 2002; Hatch et al. 2006;

Wilman et al. 2006; Edwards et al. 2009). These early studies are highly suggestive, with some objects showing rotation, offset emission and most having largely uniform [NII]/H α ratios. They also show that the most kinetically disturbed objects tend to be those which show evidence of a recent interaction. The presence of strong radio sources from the BCG at the centre of clusters (Burns et al. 1994; Hogan et al. 2015) further complicates the situation. In several clusters the intricate filamentary structures formed by the optical line emission appear to be spatially and dynamically linked to the expanding radio lobes (Conselice et al. 2001; Hatch et al. 2006), the most clear example of this is NGC 1275 in the core of the Perseus cluster.

In order to place constraints on the BCG population in general we have used the Visible Multi Object Spectrograph (VIMOS) on the Very Large Telescope (VLT) to obtain spatially resolved spectroscopy around the redshifted H α emission of 73 BCGs, spanning a range in radio power, selected to show evidence of extended line emission structures. By using this representative sample we will attempt to address three key questions pertinent to our understanding of cluster cores:

- i) *What fraction of line emitting BCGs are highly disturbed?* Previous IFU observations of just a handful of objects indicated the presence of gas at high velocity with structures and kinematics suggesting possible interactions with other cluster members (Wilman et al. 2006).
- ii) *What role does the Brightest Cluster Galaxy play in the cooling of gas from the ICM?* Most clusters of galaxies have Brightest Cluster Galaxies whose positions are strongly correlated with the cluster core (Peres et al. 1998; Sanderson et al. 2009; Hudson et al. 2010). In Hamer et al. (2012) we identified 3 objects (2 from this sample) in which the presence of ionised and molecular gas at the location of the core suggested that cooling continued despite being offset from the BCG. The systems studied in Hamer et al. (2012) had very large physical offsets but are rare.
- iii) *What role does the cold gas play in the feedback process and how is feedback fuelled?* The degree of fine tuning required for feedback to be well regulated suggests a link between the non gravitational processes in the BCG and the gas cooling from the ICM. In Hamer et al. (2014) we showed that Hydra-A contains a cold gas disc which may be responsible for moving gas into the vicinity of the AGN where it can fuel the powerful jets which are injecting mechanical energy into the ICM. If such kinematic structures are common then discs may be an important component in regulating the feedback process.

In Section 2 we present details of the observations, data reduction and analysis techniques used throughout this study. We then present the observational parameters (Section 2.4) and study the maps from individual channels of the IFU in Section 3. Section 4 then presents the analysis of the morphology and kinematics of the sample. In Section 5 we study the physical condition of the gas in the ionised nebula through its spectral properties. In Section 6 we present a discussion of our findings with regard to the questions posed earlier in this Introduction. Finally we present a summary

of our findings and the conclusions we draw from them in Section 7

Throughout this study we assume a standard cosmology with $\Omega_m = 0.27$, $\Lambda = 0.73$ and $H_o = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2 THE SAMPLE

2.1 Sample Selection and observation

This sample was drawn from the ROSAT selected Brightest Cluster Survey (BCS, Ebeling et al. 1998) and the ESO X-ray Cluster Elliptical Spectral Survey (EXCESS) which provided long slit spectra of 201 and 446 BCGs respectively. The EXCESS spectra were then used to identify clusters with optical line emission, the $H\alpha$ emission line was identified in 30% of the sample which is consistent with the findings from the BCS (Crawford et al. 1999).

Only objects with an integrated $H\alpha$ emission greater than $1 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$ were selected. The second selection criteria was that the $H\alpha$ emission showed an extent greater than two arcsec in the long slit observations to ensure only objects with line emission extended on scales greater than the expected seeing were observed.

The VIMOS instrument on the 8.2-m VLT was the primary optical IFU used throughout this work. The IFU is made up of 6400 fibres with a scale on the sky which can vary from $0.33''$ to $0.67''$ per fibre corresponding to a field of view ranging from $13'' \times 13''$ to $54'' \times 54''$. It consists of four arms which split the field of view into four identically sized quadrants. Each quadrant feeds to a separate grism which disperses the incident light onto its own CCD. There are six grisms equipped on each arm which provide overlapping coverage from 360 to 1000 nm with a spectral resolution ranging from $R \sim 200$ –2500. Due to being located in the southern hemisphere potential VIMOS targets are limited to equatorial southern objects (those with a declination of $< +20^\circ$).

The wavelength coverage allows easy access to the bright $H\alpha$ line in BCGs out to a redshift of $z = 0.25$ while the field of view is more than sufficient to contain the emission in all but the most local clusters. 78 objects matched the selection criteria of which five already had optical IFU observations. The final X-ray selected sample includes all extended $H\alpha$ bright BCGs out to a redshift $z = 0.25$ and represents an increase in the number of systems studied with IFUs of more than a factor of five. The presence of extended line emission suggests that all objects in the sample have low central entropy values (Cavagnolo et al. 2008) and are thus consistent with rapidly cooling cores. Figure 1 shows the locations of all objects from the parent samples and identifies those which were selected to be observed with VIMOS.

At high spectral resolution the kinematics of the gas can be mapped to a high degree of accuracy with each spectral resolution element (0.6 \AA) corresponding to $\sim 30 \text{ km s}^{-1}$ (at the wavelengths observed) with broad resolved lines allowing accurate centroiding to a fraction of this. The spectral resolution also allows for the independent fitting of the $H\alpha$ and [NII] lines meaning a measure of the ionisation state of the gas can be obtained from the line ratios. The wavelength coverage also allows for a detection of other important spectral features should they be present including the [SII]

doublet (allowing for a measure of the electron density), the sodium D absorption feature (which can be used to study the stellar component of the BCG) and the $H\beta$ and [OIII] lines in some objects.

VIMOS was used to obtain optical IFU spectroscopy of the 73 BCGs in the sample between October 2007 and September 2008. For each object a set of three 600s exposures were performed with a pointing dither of $1.5''$ included between each to account for bad pixels. The *HR_Orange Grism* and *GG435* filter (spectral resolution of $\sim \lambda/\Delta\lambda \sim 2650$ over a wavelength range 5250–7400 \AA) were used to observe $H\alpha(\lambda_{rest} 6562.8 \text{ \AA})$ in clusters with a redshift below 0.13. For clusters with a redshift greater than this the *HR_Red Grism* and *GG475* filter (spectral resolution of $\sim \lambda/\Delta\lambda \sim 3100$ over a wavelength range 6450–8600 \AA) were used to sample the $H\alpha$. The setup used was selected such that the spectra obtained would cover a spectral range which should include the principle lines of $H\alpha$, [OI], [NII] and [SII] assuming they are present in each objects spectrum. As this project was proposed as a bad weather backup project, the observations were taken in a range of conditions, with a median seeing of $\sim 1.5''$ but with a range ~ 0.5 – $2.6''$. Both the *HR_Orange* and the *HR_Red* modes offer a $27'' \times 27''$ field of view which is covered by 1600 fibres. Each fibre is coupled to a micro-lens to give near-continuous sky coverage with each fibre observing a region of $0.67''$ in diameter. Details of the observations are summarised in Table 1.

2.2 Data Reduction

The raw data were reduced using the VIMOS specific recipes in the ESO Recipe Execution Tool ESOREX. This package performed the basic data reduction on each data cube (a total of 12 for each object, one for each quadrant from the three pointings) including bias subtraction, flat fielding, and the wavelength and flux calibration. The wavelength calibration is achieved by comparing a separate exposure of an arc lamp which produces emission lines at known wavelengths and comparing these observations to a catalogue of known line positions. At each spatial element the lines are identified and their position compared to that in the catalogue, differences between pixels are then corrected allowing the procedure to account for instrumental distortions. The flux calibration is done using observations of a standard star (one for each quadrant of the IFU) which is compared to a standard spectrum of the star to determine the efficiency and response curves of the spectrograph. A standard extinction table is then used to calculate the flux losses due to the atmosphere at a given airmass at each wavelength, after which the scientific spectrum is multiplied by the response curve to produce the final flux calibrated spectrum. Differences in the seeing conditions and airmass of the observations and standard star are the main contributing factor affecting the flux calibrations. To minimise these effects standard stars were observed separately for each object during the observations to minimise these differences. Excluding such systematic errors the method achieves an accuracy of better than 0.5%. Full details of the data reduction pipeline and processes can be found in the VIMOS Pipeline User Manual (VLT-MAN-ESO-19500-3355).

To subtract the sky point-like objects were masked to remove any stars within the field. Such objects were defined

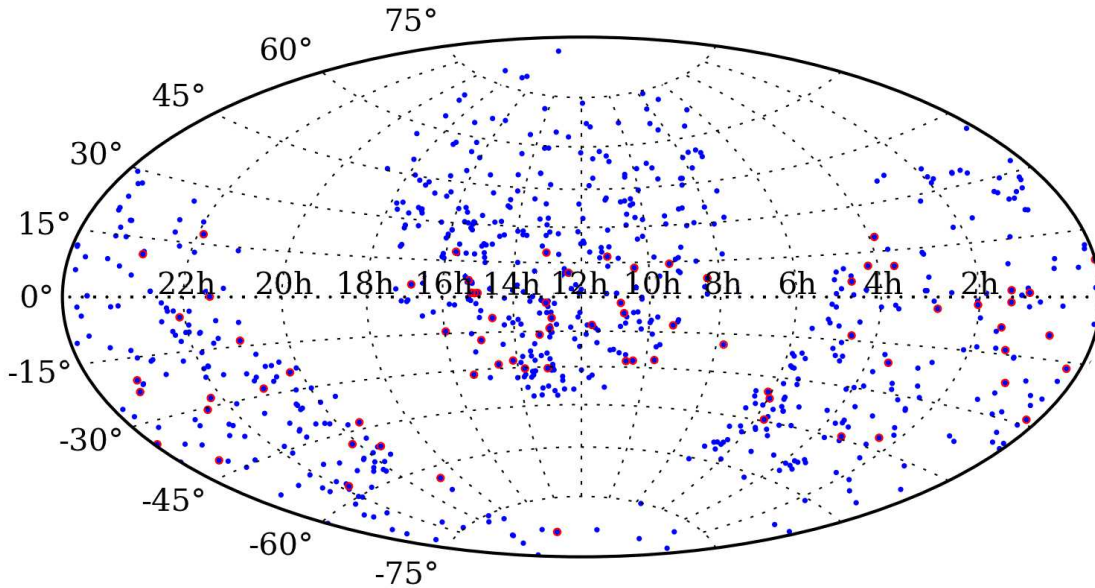


Figure 1. Here we show a plot of the sky identifying the locations of the sample. The blue points show the objects in both the BCS and EXCESS samples, the objects selected and observed with VIMOS are circled in red.

as having a roughly 2D-Gaussian intensity profile with a FWHM on the order of the seeing and were removed out to the diameter of the seeing. The BCG was then removed by masking all pixels within an isophote of half its peak intensity. Using a lower threshold (40%) had a negligible affect on the sky level (the median over all sampled spatial elements remained the same to 6 decimal places) for all but the most local objects (z less than ~ 0.01). By contrast choosing a higher threshold (60%) resulted in a median increase in the sky level across all objects of 3% suggesting that sufficient light from the BCG is being sampled to affect the sky measurement. Thus the 50% level was chosen empirically as the best compromise to allow a good sampling of sky pixels while removing the majority of the BCGs light. The small field of view means that objects other than the BCG which are not point like were rare in our sample, however some observations did contain them. These objects were removed in the same way as the BCG, using the same threshold and their peak intensity to produce the mask. The sky level for each quadrant was calculated by taking the median value of the remaining pixels at each wavelength increment. The sky spectrum was subtracted from each pixel in the four quadrants before they were combined into a single data cube.

For low redshift objects ($z \lesssim 0.02$) the BCG easily fills a single quadrant so a higher flux threshold for removing the BCG must be used which can result in a substantial over subtraction of the sky. For such objects subtraction of a straight continuum baseline prior to the sky subtraction proved a more reliable means of isolating the emission lines. The continuum baseline was calculated from the spectrum at each spatial resolution element by taking the median of the emission either side of the lines of interest. The spectral range used to calculate the median was 120 \AA in length and began 60 \AA from the emission line on either side. The continuum baseline was then produced by interpolating over the wavelength range to be fitted between the medians cal-

culated at either side of the emission lines. Once this baseline had been subtracted the sky was subtracted in the same manner as the higher redshift objects before the lines are fit. The drawbacks of this method however are that it is only valid for a small spectral range so only one line complex can be studied in a given reduction and it also eliminates any information carried in the objects continuum emission.

Finally the three exposures were median combined for each pointing in order to eliminate cosmic rays. The exception to this was RXCJ2014.8-2430 which only had two pointings so the mean of the two observations was used to produce the final cube. The final mosaics provided cubes with a $\sim 30'' \times 30''$ field of view once the dithering offsets are taken into account.

2.3 IFU Data Cube analysis

We took the IFU data cubes produced by the reduction and fit spectral models to the key lines at each resolution element. The primary diagnostic line structure was the $H\alpha$ ($\lambda_{rest} 6562.8 \text{ \AA}$) and $[NII]$ ($\lambda_{rest} 6548.1/6583 \text{ \AA}$) complex.

To fit these lines we extracted a region of the spectrum in each resolution element (lenslet) that was 240 spectral resolution elements in length ($\sim 144 \text{ \AA}$) centred at the wavelength of the $H\alpha$ line at the cluster redshift. We then fitted Gaussian emission-line profiles to these extracted regions using a χ^2 minimisation procedure. To the $H\alpha$ - $[NII]$ triplet of lines we fit three separate Gaussian profiles on top of a flat baseline to account for variations in the average continuum flux and compare this to just a flat continuum baseline fit (Figure 2). Each of the three emission line had the form

$$F(\lambda) = \frac{F_0}{2\pi\sigma} \times e^{\frac{-0.5(\lambda-\lambda_c)^2}{\sigma^2}} \quad (1)$$

with F_0 the total flux in the line, σ the velocity dispersion of the line ($\text{FWHM}/2(2\ln 2)^{0.5}$), λ the wavelength at a given position and λ_z the wavelength of the line corrected for redshift. The model spectrum contains 3 such lines with fixed (at redshift 0) separations, the $\text{H}\alpha$ -[NII] $_{\lambda 6583/\lambda 6548.1}$ triplet. Each line was produced independently and the resulting spectral models were added together on top of a flat baseline to produce the final spectral model.

The fitting routine was given an array of first guess variables which were then used to produce the fit. For the $\text{H}\alpha$ -[NII] triplet these variables were as follows:-

- (i) The redshift of the emission. The redshift was used to determine the value of λ_z in Equation 1 such that $\lambda_z = \lambda_0 \times (1+z)$. Where λ_0 is the rest wavelength of the line and z is the redshift. The first guess for the redshift at each position was set to the systemic redshift of the BCG.
- (ii) The total flux in the $\text{H}\alpha$ line. This was used as F_0 for the $\text{H}\alpha$ line in Equation 1. The first guess was set to the average flux at the wavelength of the line across the whole cube.
- (iii) The total flux in the [NII] $_{\lambda 6583}$ line. This was used as F_0 for the [NII] $_{\lambda 6583}$ line in Equation 1, F_0 for the [NII] $_{\lambda 6548.1}$ line was fixed at one third of this value. The first guess was set to be equal to the $\text{H}\alpha$ flux.
- (iv) The FWHM of the line in angstroms. This was used to calculate σ in Equation 1 such that $\sigma = \text{FWHM}/2(2\ln 2)^{0.5}$. The first guess set to the width of a skyline, $\sim 4.7 \text{ \AA}$.
- (v) The baseline flux level. This was not used to produce any of the three line models but was added to the total model spectrum after they had been produced to account for variations in the median level of the continuum between lenslets.

The redshift and width of the [NII] and $\text{H}\alpha$ lines were set to be the same to reduce the number of free parameters in the model. After producing the first guess the redshift, $\text{H}\alpha$ flux, [NII] $_{\lambda 6583}$ flux, linewidth and continuum baseline level were allowed to vary using a least squared fitting routine (Moré & Wright 1993) until the model minimised to best fit the data. During the minimisation the parameters were constrained to ensure only the emission lines of interest were recovered and reduce the processing time required to achieve a fit. We constrained the $\text{H}\alpha$ and [NII] $_{\lambda 6583}$ flux to positive values (emission) and the sigma linewidth to $< 1000 \text{ km s}^{-1}$ ($\text{FWHM} \sim 2355 \text{ km s}^{-1}$). These profiles were initially fitted to the spectrum from each $0.67''$ lenslet and adaptively binned to $\sim 2.0''$ (3×3 lenslets) in regions with lower $\text{H}\alpha$ flux, as such the fits to the low surface brightness emission have a lower spatial resolution than the brightest regions.

We accepted these fits as representing the presence of spectral lines in the data when they provided an improvement in signal to noise of 7 (7 times the standard deviation, or 7σ) when compared to a continuum baseline only fit. When an acceptable fit was found the parameters of the best fitting model were stored and the 1σ errors were calculated by varying each parameter slightly from the best fit and allowing all others to re-minimise. For each pixel where an acceptable fit was found we attempted to fit a second model which included an additional line component with a redshift, intensity and width independent of the initial fit. This extra component was allowed to become much broader than the initial fit (it was left unconstrained) and was included as a means to test the spectra for additional velocity components and broad line features. This model was then ac-

cepted when it provided an improvement in signal to noise of 15 (15 times the standard deviation, or 15σ) when compared to a continuum baseline only fit. For pixels where a second component was found to be significant a model including a second component for all lines was fitted to the data and the errors were recalculated to account for the new model. This process detected only two objects with multiple velocity components and eight with broad components in their core regions. As multiple velocity components do not make up a significant fraction of the sample we choose not to address the nature and origin of the additional components in this paper. The significance of these models was tested on a number of objects using an F-test of additional parameters. The basic $\text{H}\alpha$ -[NII] line model when compared to a baseline only model was found to always have a significance level of $\alpha < 0.001$ when required to fit a signal to noise of 7σ . Likewise the threshold of a signal to noise of 15σ was selected so as the additional line components had a significance level of $\alpha < 0.001$ when compared to the standard $\text{H}\alpha$ -[NII] line model.

The parameters of the fits were then used to produce maps of the line flux, line-of-sight velocity and full width at half-maximum (FWHM) deconvolved for instrumental resolution. The velocity zero point was determined as the median of the measured velocities after removing the highest 2.5% and lowest 2.5% of the values. A retrospective clipping method was used to remove pixels in which the fitting routine had returned parameters with non physical properties (velocities in excess of $\pm 1000 \text{ km s}^{-1}$ relative to the mean which are unlikely to be associated with the BCG). Continuum images of the region covered by the observations were produced by taking the median from each lenslet over a region of the spectra (here after referred to as collapsing the cube) containing no emission lines or sky line residuals. These were produced too allow comparison between the position of continuum objects and the structure of the $\text{H}\alpha$ emission to identify possible offset peaks (Hamer et al. 2012) or interactions with other cluster members.

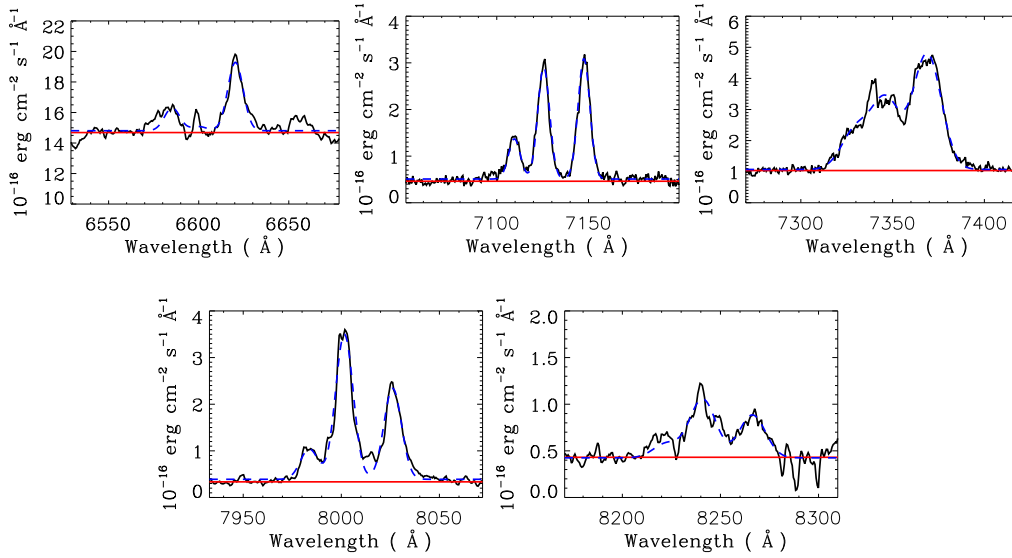


Figure 2. Here we show several spectra showing the $H\alpha$ and $[NII]$ lines which represent the sample. The spectra are of the central region of five objects and the wavelength coverage is the same as the region used by the fitting routine. From left to right, top to bottom the redshifts of the objects are ≈ 0 (NGC 5846), ≈ 0.08 (Abell 478), ≈ 0.13 (Abell 1348), ≈ 0.21 (Abell 3017) and ≈ 0.26 (Abell 3444) so are representative of the whole sample. The red solid line shows the flat continuum baseline fit and the blue dashed line shows the best fit spectral model. The variation of flux, $[NII]/H\alpha$ ratio and line width are apparent from these plots.

2.4 Observational Parameters

Table 1 outlines the key observational parameters of each object in the sample. The redshift quoted in Table 1 is the median redshift of the $H\alpha$ line in the object. Most of the objects in the sample had all pointings performed on the same day, or at least in similar conditions. However, RXCJ2014.8-2430 had one of its pointings performed in very poor conditions. As such we have only used two of the pointings to make the final cube. RXCJ2101.8-2802 had pointings which were taken on different days, with 4 pointings in total, all of which were in consistent conditions and showed no other problems in the spectra so we choose to use all 4 pointings, giving this object a slightly longer integration time. We calculated the mean seeing of all observations to be ~ 1.5 arc-sec. Before the full fitting routine was run on an object its spectrum was studied to identify the presence of additional emission lines (lines other than $H\alpha$ and $[NII]$), these additional lines are listed in Table. 1. We base our classifications on the x-ray properties of each object (x-ray luminosity and where available x-ray temperature) using the classification of Bahcall (1999) (Groups $< L_x = 10^{43} \text{ erg s}^{-1} < \text{Clusters}$ and Groups $< T_x = 2 \text{ keV} < \text{Clusters}$). By this definition our sample consists of 61 brightest cluster galaxies and 12 brightest group galaxies.

TABLE 1.
THE OBSERVATIONAL PARAMETERS LISTED FOR EACH OBJECT IN THE VIMOS SAMPLE

Cluster	Redshift	Grating	Exposure (s)	Mean Seeing(")	Lines Detected	RA	Dec	L_x (10^{43} erg s $^{-1}$)	T_x (keV)	Classification
Abell 1060	0.01263	HR-orange	3×600	0.79		10:36:42.78	-27:31:39.66	3.34	3.2	Cluster
Abell 1084	0.13301	HR-red	3×600	0.68	[OI]	10:44:32.69	-07:04:05.21	48.99	7.3	Cluster
Abell 11	0.14910	HR-red	3×600	0.77	[OI], [SII]	00:12:33.58	-16:28:05.17	18.2	—	Cluster
Abell 1111	0.16518	HR-red	3×600	0.85	[OI], [SII]	10:50:36.26	-02:36:13.37	27.30	—	Cluster
Abell 1204	0.17057	HR-red	3×600	0.81	[OI]	11:13:20.08	17:35:37.88	72.6	3.78	Cluster
Abell 133	0.05665	HR-orange	3×600	0.93	[OI], [SII]	01:02:41.49	-21:52:48.62	16.17	3.71	Cluster
Abell 1348	0.11985	HR-orange	3×600	0.87	H β , [OIII], [NI], [OI]	11:41:24.12	-12:16:35.90	21.09	5.5	Cluster
....										

Table 1. The observational parameters listed for each object in the VIMOS sample. The first column states the cluster designation as used throughout this paper. The redshift is the median redshift acquired from the fits to the H α -[NII] complex, the median is calculated from the fits after dropping the top and bottom 2.5% of the redshifts. The grating is the VIMOS grating used for the observations and the exposure lists the total integration time on target. The seeing is calculated as the mean of the DIMM seeing across all exposures for each object. Lines detected lists all lines apart from the H α and [NII] lines that were detected by visual inspection, the H α -abs line is included for systems which show some evidence of H α absorption. The right ascension and declination are measured at the centre of the field of view for each observation. X-ray luminosities were taken from the surveys of Böhringer et al. (2004); Ebeling et al. (1998, 2000) and temperatures are taken from Ebeling et al. (1996, 1998, 2000); Cavagnolo et al. (2009). The classifications are based on the x-ray properties following Bahcall (1999). The full table can be found in Appendix ??

3 THE RAW DATA: CHANNEL MAPS

We begin by presenting channel maps of several representative objects produced from the fully analysed data cubes which show the H α emission at different velocities. To produce the channel maps we baseline subtracted each spatial resolution element by fitting a straight line to the spectrum. This was then subtracted to remove the bulk of the continuum emission from the cube. Channels at greater than ± 1000 km s $^{-1}$ relative to the redshift of the BCG were then used to determine the average continuum subtracted channel noise of the observation, these channels were then discarded. The remaining channels contained the H α emission from the cluster core. For channels where the continuum subtracted emission exceeds 7 times the average continuum subtracted noise a flux map was produced to show the location of the H α emission in that channel. Each map gives the velocity of its channel with respect to the median velocity of the H α emission for that object. For reference the centre of the continuum emission from the BCG for each object is marked on its respective maps with a blue cross. This position was determined by fitting 2-dimensional Gaussian profiles to a collapsed continuum image.

We first show the channel map of Hydra-A (Figure 3) which shows evidence for a rapidly rotating disc within its core (Hamer et al. 2014). This is apparent from the channel maps which show a shift in the position of the bright component of the H α emission as the velocity changes. Additionally we see that the filaments extending to the north and south also show a velocity shift, with the northern filament being slightly blueward of the core and the southern filament shifted towards the red. This velocity shift is not clearly apparent in the velocity map displayed in Hamer et al. (2014) as the shift is small relative to the total velocity shift in the system. This kind of velocity structure, with a shift in the velocity of the bulk of the emission across the BCG, is common amongst the objects in the sample clearly visible in the channel maps of 31/73 clusters. The velocity shift in most

is not as extreme as that seen in Hydra-A, however this is not surprising given we are viewing the disc in Hydra-A at close to edge on (Hamer et al. 2014).

NGC 5044 has narrow band imaging that shows an extensive network of H α filaments (Werner et al. 2014) similar to those seen in the archetype cooling flow NGC 1275. Our VIMOS observations do not resolve the individual filaments but do indicate diffuse extended emission to the south in the region where the filaments are found. We see this emission extending to the south in channels close to the BCGs central velocity (Figure 3). The kinematics correlate well with those of the molecular CO(2-1) emission seen by David et al. (2014), in the centre the line profile is broad (FWHM ~ 400 – 500 km s $^{-1}$) and double peaked suggesting 2 velocity components separated by ~ 300 km s $^{-1}$ are present. Multiple velocity components such as this are only apparent in the channel maps of NGC 5044. However they have been observed in IFU observations of other systems (e.g. Abell 1664 Wilman et al. 2006, 2009) warranting a more detailed search for their presence. The channel maps of NGC 5044 show that the brightest region of the emission can be found at different positions within the BCG at different velocities. It is slightly extended towards the northwest at velocities blueward of the BCG (< -200 km s $^{-1}$), runs from the northeast to southwest of the BCG at ~ -120 – 0 km s $^{-1}$ and extends slightly to the southeast when redward of the BCG (> 150 km s $^{-1}$) suggesting that it has no ordered velocity structure. This redshifted region matches the position and velocity relative to the centre of the system of GMA 18 identified in David et al. (2014). Low surface brightness extended regions of H α emission similar to those seen in NGC 5044 are seen in 9/73 objects in our sample suggesting that while other filamentary structures may be present in our sample, they are not common. The channel maps for all objects can be seen in Appendix ??.

It is clear from the channel maps that there is a wide variety of morphologies of the H α emitting gas between ob-

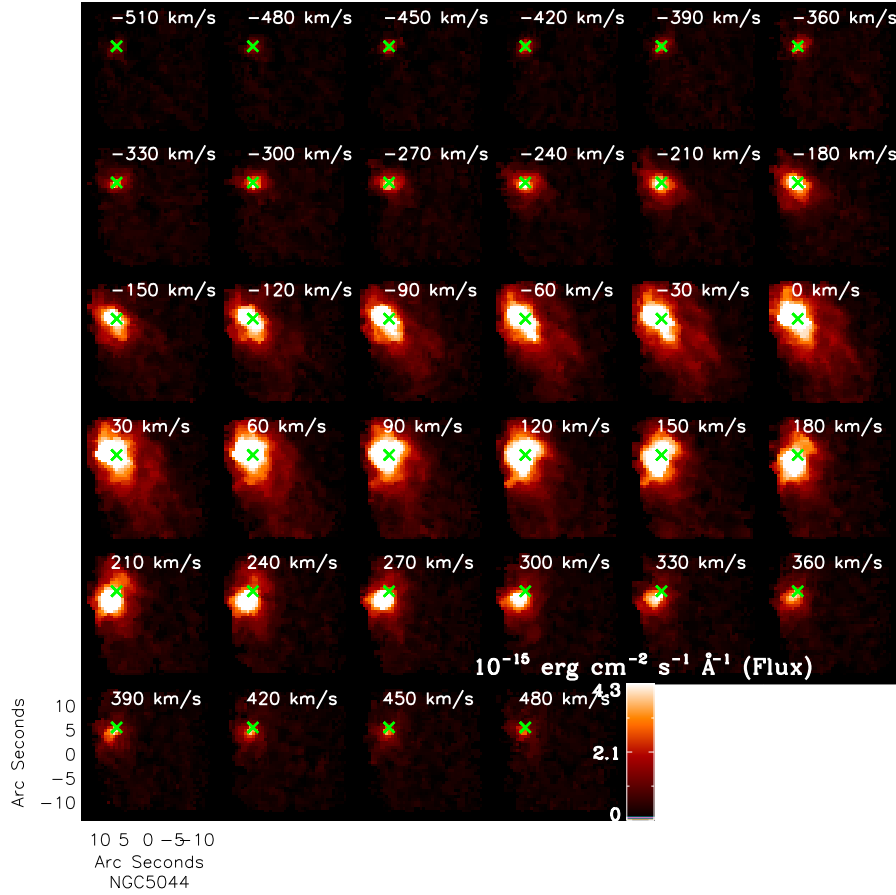
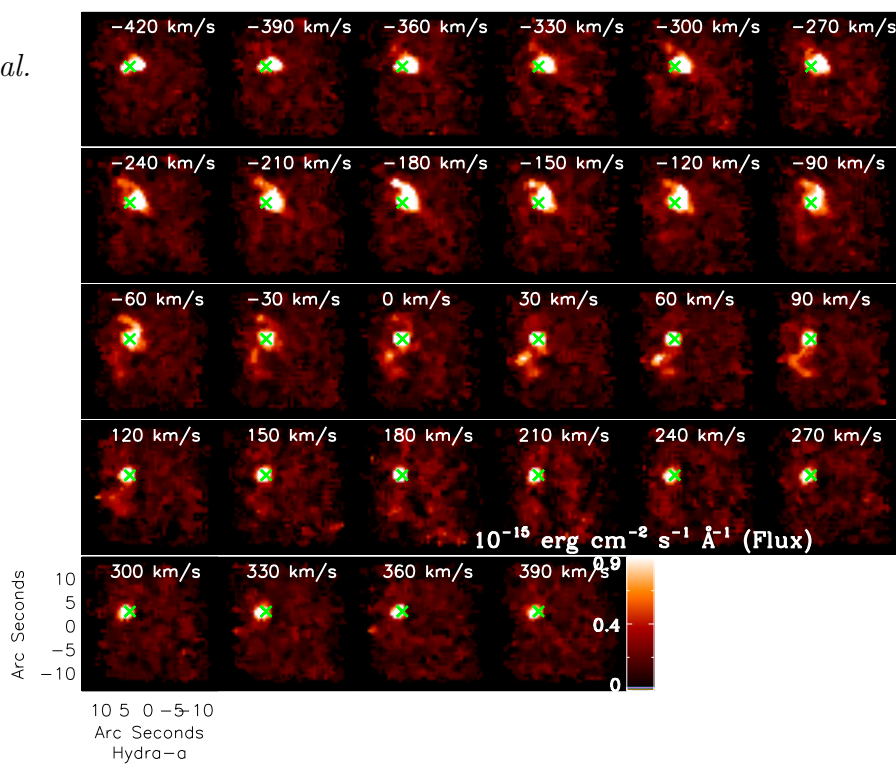


Figure 3. *Top* - Here we display continuum subtracted maps of the H α emission in Hydra-A for each channel containing emission. Each map gives the velocity of its channel with respect to the median velocity of the H α emission. The green cross marks the centre of the continuum emission from the BCG. The maps clearly show a shift in the centroid of the H α emission as the velocity changes which follows the disc within this object. *Bottom* - The channel maps of NGC 5044 showing the H α emission in velocity slices. The blue cross marks the centre of the continuum emission from the BCG. These maps show a low surface brightness region of extended diffuse emission to the south where filaments are seen in narrow band images. The core region of the emission shows a chaotic variation in structure between velocity channels suggesting the kinematics in the core are quite complex and disturbed. Channel maps for every object can be found in Appendix ??

jects. Some objects appear very compact from their $H\alpha$ emission while others are very extended. Some of this results naturally from the variation in redshift of the objects within the sample however, there remains a large variation when we consider the physical extent of the objects (See Table 2).

In addition to their varied spatial morphology, many systems show a significant, coherent velocity gradient across the BCG (31/73). This suggests the presence of an ordered velocity field within the gas which for one object (Hydra-A) we know traces a rotating disc of gas that aligns with the radio and X-ray structure in the system (Hamer et al. 2014). The presence of such a velocity structure in a large fraction of BCGs would have important implications for our understanding of cluster cores.

4 THE STRUCTURE AND KINEMATICS OF THE GAS

In order to parameterise the data and make it easier to analyse we fit Gaussian emission line profiles to the $H\alpha$ and [NII] lines in each lenslet (see Section 2.3). We now study the maps of these parameterised fits to the $H\alpha$ and [NII] triplet for all 73 objects in our sample. Five panels were produced for each object showing the continuum emission, the $H\alpha$ flux map, the ratio of the [NII] to $H\alpha$, the line of sight velocity field and the FWHM. It should be noted that in some systems where the continuum is very high and $H\alpha$ flux is low not all of the $H\alpha$ is recovered due to the underlying absorption feature caused by $H\alpha$ in stellar atmospheres. In such cases the $H\alpha$ line often appears very weak or even completely absent from some parts of the cube where it is otherwise expected to be (such as regions where the other ionised lines are strong). Systems with this problem were typically of low redshift and thus filled much of the field of view making producing an empirical model of the absorption from a bright region of the galaxy devoid of line emission difficult or impossible.

Rather than trying to model out the absorption as a component of the fitting routine (which would have reduced the degrees of freedom for these objects) we fitted the spectrum as normal. In these cases the spectral fit was primarily constrained by the bright [NII]($\lambda_{\text{rest}} 6583 \text{ \AA}$) line which could introduce some bias as the two lines may be tracing different excitation mechanisms (AGN processes typically produce more [NII] for a given $H\alpha$ than star formation for example). However, in all such cases the [NII]/ $H\alpha$ ratio is high so [NII] provides a higher significance tracer of the gas dynamics without the degeneracy of additional model parameters needed to model out the absorption. In such cases the maps show the [NII] flux rather than the $H\alpha$.

To determine the effect of the bias introduced by this approach we produced maps constrained only by the [NII] for systems where $H\alpha$ and [NII] were equal in strength and compared them to the $H\alpha$ constrained maps. We found no significant difference between the maps suggesting that the [NII] emission is tracing the same morphology and kinematics as the $H\alpha$, at least in the case of the galaxies we could perform this test on. We adopt the assumption that this concordance extends to objects for which $H\alpha$ emission was too weak and the stellar absorption component dominates to produce reliable maps constrained by $H\alpha$. Figure 4 shows

maps of the parameterised fits for a sub sample of objects, the maps of the full sample are available in Appendix ??.

4.1 $H\alpha$ morphology

From the maps shown in Figure 4 we can calculate the extent of the line emitting region for each object. We define the major axis to be the maximum length of continuous emission (at greater than 7σ) along a line through the peak of the $H\alpha$ emission (throughout this paper we use extent and major axis interchangeably). The minor axis is then the length of continuous emission perpendicular to the major axis (also through the peak of the $H\alpha$). This definition assumes that the distribution of $H\alpha$ emission is well approximated by an ellipse. This is clearly not the case for all objects in the sample though it is a good approximation to regions of uniform $H\alpha$ distribution in most systems. We note however, that it is also important to characterise the overall distribution of the emission for those objects which are extended.

As can be seen from Figure 4 the $H\alpha$ morphology varies greatly from system to system. We categorise these morphologies into five distinct groups. It is important to note that the definitions are not all mutually exclusive, as such it is possible that a single object may exhibit a morphology which is consistent with two or more of the definitions given below.

Compact objects Compact objects are defined as objects in which the minor extent of the $H\alpha$ emission above a significance of 7σ is less than twice the mean of the DIMM seeing during the observation of that object. The selection criteria of the VIMOS sample required that every object has an extent of more than $2''$ in the FORS spectra. As such very few objects should fall into this definition however we include it to account for observations taken in very poor seeing conditions. It is important to note that this definition of compact makes no consideration of the physical scale of the objects. However, it is an important classification as any apparent spatial variations along the minor axis seen for such objects cannot be believed. Objects that are given this classification are also classified according to the structure seen along their major axis unless this too is less than twice the seeing.

Plumes An object is classified as having a plume when the $H\alpha$ emission shows a clear extent in one preferential direction which is not shared by the continuum emission. Typically with this type of object the bright central region of the $H\alpha$ emission also shares the shape of the larger scale plume which results in the peak of the $H\alpha$ emission being slightly offset from the peak in the continuum. This offset is never a large physical separation ($\lesssim 7 \text{ kpc}$) but is an important parameter to distinguish plumes from quiescent objects in which $H\alpha$ emission is only detected to one side of the BCG, the latter showing no offset larger than the mean seeing.

Offset objects Objects classified as offset show a $H\alpha$ morphology in which either all of the emission is offset from the peak of the continuum or that have a second $H\alpha$ peak which contains the majority ($> 50\%$) of the $H\alpha$ flux that is substantially offset from the BCG ($> 8 \text{ kpc}$, the maximum physical extent of the seeing in any object from the sample). Although we classify both of these definitions as

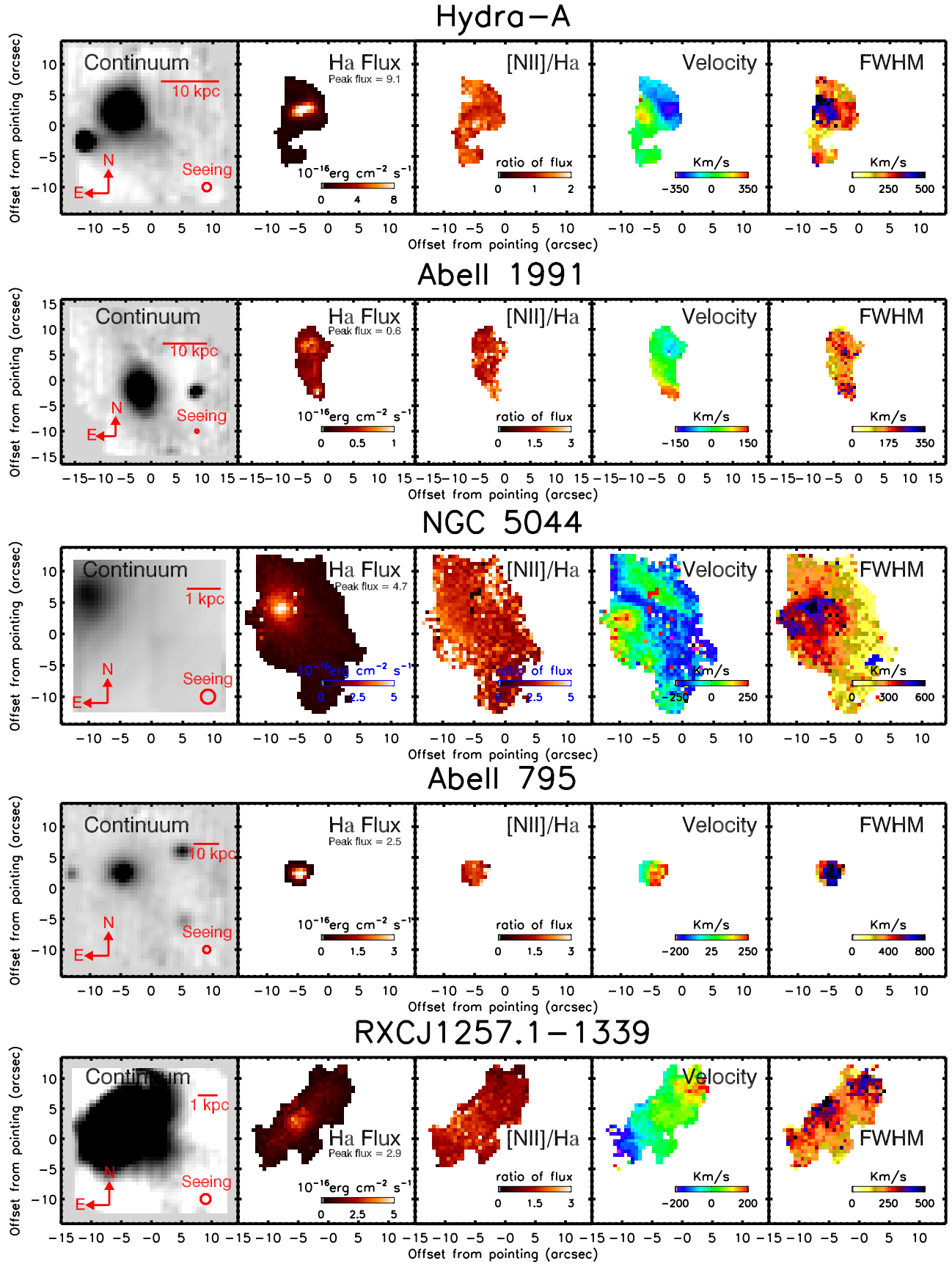


Figure 4. Maps of the spectral fits to the VIMOS data cubes for the five objects from the sample. From left to right the panels are: 1) A continuum image created by collapsing the cube over a wavelength range free of emission and sky lines, 2) The $H\alpha$ flux map, 3) A map showing the ratio of [NII] to $H\alpha$ flux, 4) the line of sight velocity profile of the $H\alpha$ and [NII] emission, 5) the Full Width at Half Maximum (FWHM) of the emission line deconvolved for instrumental resolution.

offset they show distinctly different $H\alpha$ morphologies. The objects which are completely offset from the centre of the BCG appear to be otherwise quite similar to quiescent objects in their morphology. The latter kind however, appear morphologically more consistent with plumes with a clear extent of the low surface brightness $H\alpha$ emission extending in a single preferential direction. However, there are two key differences to plumes. First, the extent seen in the low surface brightness emission is not shared by a similar structure in the bright emission as is the case with plumes. Secondly, those which are classified as offset have a second peak in the $H\alpha$ emission which is positioned at the end of the extended region of low surface brightness emission, away from the BCG. For a more detailed analysis of the nature of objects with this kind of offset emission see Hamer et al. (2012).

Quiescent objects Objects classified as quiescent have $H\alpha$ emission which is extended beyond twice the seeing and shows a simple elliptical and centrally concentrated morphology. The peak of the $H\alpha$ emission must lie within the mean seeing of the peak in the continuum and the $H\alpha$ flux must fall off uniformly as the continuum does (though not necessarily at the same rate as the continuum). Objects with a $H\alpha$ morphology of this type are likely to be very relaxed and have had no major events occur recently enough to disturb the gas.

Disturbed objects The final morphological classification we present are disturbed objects. These are objects where the $H\alpha$ flux typically peaks close to the peak of the continuum but have lower surface brightness $H\alpha$ emission which is extended and forms non-uniform structures around the BCG. This classification of object essentially includes all objects which are not classified under one of the previous four categories. The extended and non-uniform nature of the $H\alpha$ emission in this class of objects suggests that some event has disturbed the cold gas in the core of the cluster causing it to flow away from a uniform distribution. Another possibility is that this emission is tracing the gas being accreted by the system. Some of the objects which fall into this category are systems which are well known to have active AGN which can be seen to be disturbing the ICM (for example Abell 2052 Blanton et al. 2011). This suggests the possibility that objects which fall into this classification are currently undergoing feedback processes which has resulted in the disturbance of the cold gas reservoir.

The classification assigned to each object in the sample can be found in Table 2. Some care must be taken when considering these classifications (in particular the distinction between quiescent and disturbed objects) as the surface brightness limits of our objects ranges considerably, from $0.38 \times 10^{34} \text{ erg s}^{-1} \text{ kpc}^{-2}$ to $1.61 \times 10^{37} \text{ erg s}^{-1} \text{ kpc}^{-2}$. We note however, that these are extreme examples and the sample has a median surface brightness limit of $3.97 \times 10^{35} \text{ erg s}^{-1} \text{ kpc}^{-2}$ with a median absolute deviation of $3.19 \times 10^{35} \text{ erg s}^{-1} \text{ kpc}^{-2}$. A histogram of the surface brightness limits grouped by classification can be seen in Figure 5. The disturbed objects (listed in Table 2) sample the full range of surface brightness limits ranging from $1.90 \times 10^{34} \text{ erg s}^{-1} \text{ kpc}^{-2}$ to $1.64 \times 10^{37} \text{ erg s}^{-1} \text{ kpc}^{-2}$. However, the median surface brightness limit of the disturbed objects is $1.03 \times 10^{36} \text{ erg s}^{-1} \text{ kpc}^{-2}$ with a median absolute deviation of $8.22 \times 10^{35} \text{ erg s}^{-1} \text{ kpc}^{-2}$ which does suggest that

objects classified as disturbed typically have higher surface brightness limits than average for the sample but is consistent with the median of the full sample within the errors. While the small sample size of disturbed objects (just 13) must be considered when interpreting these values the large range of surface brightness limits measured for the sample suggest that the surface brightness limit has only a small effect on our ability to distinguish an objects morphology. The seeing of our observations may also play a role in our ability to classify objects correctly, we test this in Figure 6 by plotting the seeing of the observations against the surface brightness limit for the object. No obvious trends are visible on this plot suggesting that seeing does not have a major impact on our ability to identify an objects morphology. However, we do note that all offset objects were observed in relatively good seeing conditions. This is not unexpected as a poorer seeing will limit our ability to accurately determine the position of the $H\alpha$ peak and thus identify the object as offset if the separation is small.

High resolution narrow-band imaging of $H\alpha$ in cluster cores has shown the presence of filamentary structures surrounding the BCG. While a few objects in our sample do show narrow filaments of emission which extend out into the ICM we do not resolve these structures in the majority of cases. Most likely this is a result of the poorer seeing of our observations smoothing the filamentary structures to such an extent that they appear blended as a continuous region of low surface brightness emission extended away from

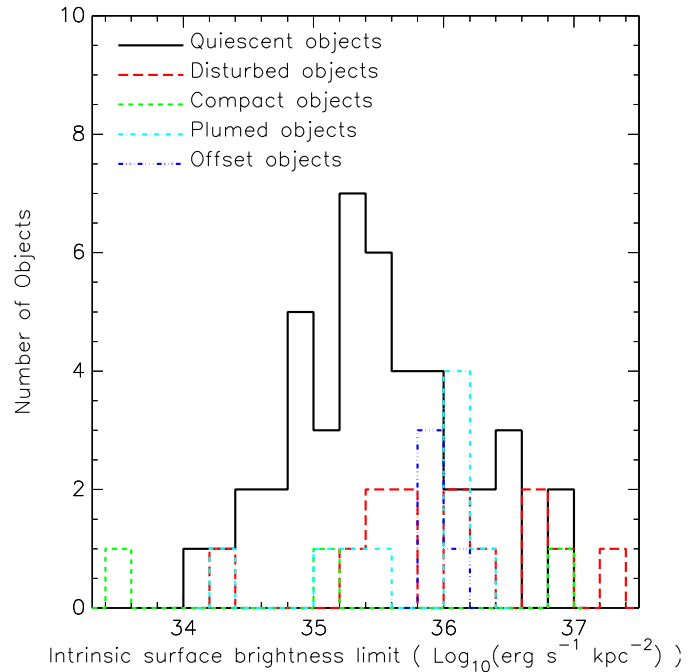


Figure 5. A histogram showing the distribution of intrinsic surface brightness limits for each classification of object. We note that, with the exception of offset objects, all classifications sample the majority of the range of surface brightness limits. Interestingly, the plot does show that 11 out of 14 objects with a surface brightness limit of below $10^{35} \text{ erg s}^{-1} \text{ kpc}^{-2}$ are classified as quiescent suggesting that the higher surface brightness features associated with more complex morphologies can be identified in all systems.

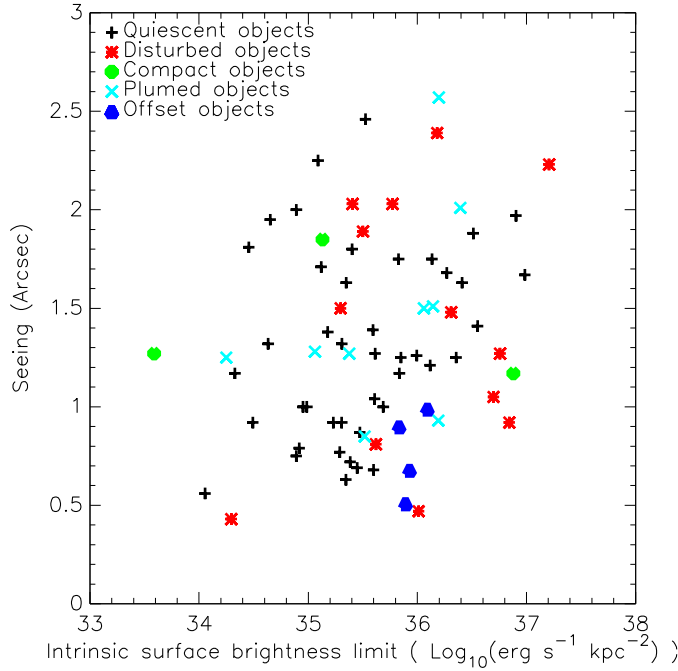


Figure 6. A plot showing the mean seeing of the observations for each object against that objects intrinsic surface brightness limit. Each object is given a different symbol depending on its classification. If the seeing of our observations played a major role in our ability to distinguish the morphology of an object we would expect to see trends on this plot. However, with the exception of offset objects, all others fill the parameter space with no obvious trends. The offset objects are only seen at good seeing, $\sim 1''$ or below.

the BCG. A good example of this effect is NGC 5044 which is known to have filaments of $H\alpha$ emission extended to the south-west of the BCG (David et al. 2011). Comparing to our $H\alpha$ flux map shown in Figure 4 it is possible to see low surface brightness emission extended to the south-west over similar distances to the filaments, however, the details of the filamentary structures are not visible in the VIMOS observation. In the few cases where we do see structures which appear similar to filaments we note that their apparent thickness is substantially higher than those observed with narrow-band imaging. While this may be an effect of the seeing it is impossible to tell without high resolution $H\alpha$ imaging of these systems.

4.1.1 Offset Emission

In Hamer et al. (2012) we report on the nature of three objects (two from this sample) which show a significant component (> 50 per cent) of their line emission significantly offset from the centre of the BCG. Figure 7 (top) shows a histogram plot of the apparent visible offset between the BCG and the majority of the clusters line emission. It can be seen that almost all clusters show a small offset of less than $4''$ and most well below the seeing limit (~ 1.5 arcsec). Only one object (Abell 1991) shows an offset of greater than this which makes it clearly stand out amongst our sample. In the middle panel of Figure 7 we present the physical offset

of the BCG from the majority of the line emission calculated at the redshift of the host cluster.

We draw three conclusions from this plot, the first is that the majority of the objects show a small physical offset as is expected. Second, we note that there are now two objects showing a significant offset, one at ~ 10 kpc and one at ~ 12.5 kpc. These objects are Abell 1991 and Abell 3444 which were studied in Hamer et al. (2012). Additionally, there is a cluster of objects with an offset of ~ 6 kpc in Figure 7. Studying the maps of these clusters, we note they are mostly plumes rather than separate offset emission. However, one of these objects is Abell 2566 which is offset and shows no line emission at the location of the continuum peak. In this respect it is similar to the Ophiuchus cluster identified by Edwards et al. (2009).

Hamer et al. (2012) postulated that the apparent differences between Ophiuchus and Abell 1991/ 3444 may have been the effect of viewing the offsets inclined at different an-

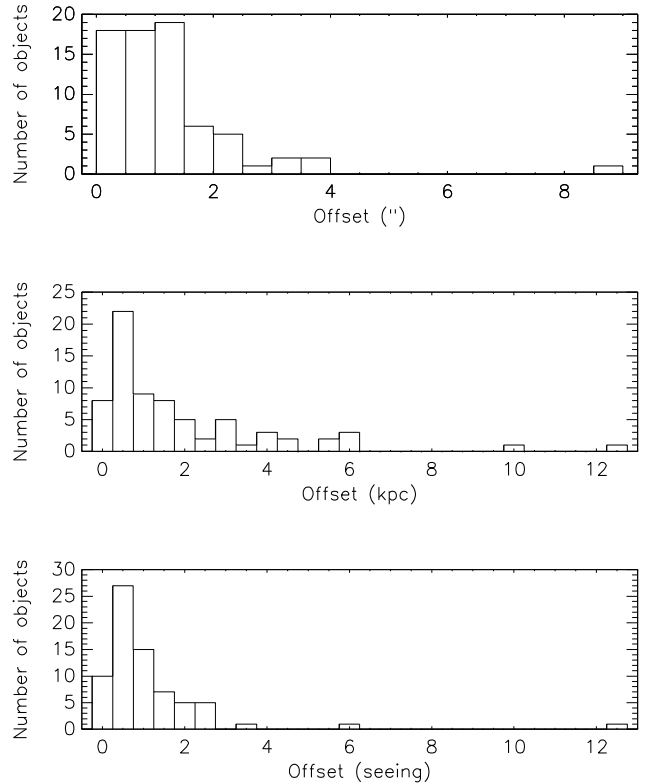


Figure 7. The top plot (a) shows the apparent offset between the continuum emission of the BCG and the centroid peak of the $H\alpha$ emission for the full sample of objects. Almost all objects show small apparent offsets of less than $4''$. Only one object shows a significant apparent offset, this is Abell 1991. The middle plot (b) shows the physical observed offset of the BCG and the majority of the line emission. Here we see a more spread out distribution of objects but with the majority still showing small offsets. The bottom plot (c) shows the offset as measured in terms of seeing units (offset in arcsec / mean seeing). We note that most objects are offset by less than 1 seeing unit confirming that there is no real offset in these objects. there are also two objects which show a very large offsets at 6 and 12.5 seeing units. These are Abell 3444 and Abell 1991 respectively and clearly indicates that the offset in these objects is real.

TABLE 2.
 THE DERIVED PARAMETERS LISTED FOR EACH OBJECT IN THE VIMOS SAMPLE

Cluster	Velocity km s ⁻¹	Extent kpc(arcsec)	$L_{H\alpha}$ 10 ⁴⁰ erg s ⁻¹	FWHM km s ⁻¹	[NII]/H α	Morphology
Abell 1060	68 ± 11	1.22(4.73)	0.043 ± 0.0086	107 ± 17	0.7 ± 0.3	Quiescent
Abell 1084	269 ± 30	11.23(4.75)	11 ± 1.8	275 ± 108	0.9 ± 0.5	Quiescent
Abell 11	247 ± 15	29.58(11.37)	70 ± 6.5	340 ± 61	0.9 ± 0.5	Quiescent
Abell 1111	478 ± 22	21.94(7.75)	33 ± 3.9	297 ± 75	1.1 ± 0.5	Plume
Abell 1204	490 ± 29	26.32(7.15)	59 ± 5.4	514 ± 355	1.7 ± 0.6	Disturbed
Abell 133	210 ± 32	7.18(6.53)	1.2 ± 0.16	308 ± 14	1.7 ± 0.2	Plume
Abell 1348	755 ± 40	18.05(8.35)	46 ± 6.3	513 ± 153	1.6 ± 0.5	Quiescent
Abell 1663	225 ± 45	13.71(8.66)	1.7 ± 0.29	526 ± 120	3.4 ± 0.4	Quiescent/Compact
Abell 1668	302 ± 22	10.90(8.92)	4.7 ± 0.45	339 ± 28	2.2 ± 0.7	Quiescent
Abell 194	85 ± 7	1.49(4.03)	0.50 ± 0.071	211 ± 122	0.8 ± 0.23	Compact
Abell 1991	124 ± 10	16.41(14.38)	4.0 ± 0.35	209 ± 64	1.5 ± 0.3	Offset
Abell 2052	268 ± 18	6.44(9.41)	1.8 ± 0.18	454 ± 105	2.2 ± 1.1	Disturbed
Abell 2390	490 ± 29	25.99(7.06)	109 ± 15	422 ± 52	0.95 ± 0.48	Quiescent
Abell 2415	186 ± 37	11.25(10.08)	4.5 ± 0.42	275 ± 180	1.5 ± 0.55	Plume
Abell 2495	201 ± 16	8.88(5.93)	0.56 ± 0.090	235 ± 38	1.5 ± 0.074	Quiescent
Abell 2566	325 ± 22	10.18(6.5)	9.7 ± 1.2	172 ± 74	0.91 ± 0.56	Offset
Abell 2580	203 ± 24	9.40(5.71)	2.4 ± 0.48	340 ± 157	0.75 ± 0.55	Compact/Plume
Abell 2734	210 ± 20	6.09(5.09)	1.1 ± 0.27	551 ± 105	2.5 ± 1.0	Quiescent
Abell 291	294 ± 70	23.14(7.12)	480 ± 31	626 ± 371	0.99 ± 0.55	Disturbed
Abell 3017	390 ± 31	25.03(7.06)	110 ± 17	366 ± 106	0.62 ± 0.33	Quiescent
Abell 3112	225 ± 19	8.94(6.20)	7.1 ± 1.3	893 ± 256	1.9 ± 1.2	Quiescent
Abell 3378	106 ± 15	20.63(8.30)	6.1 ± 0.95	422 ± 170	1.1 ± 0.39	Quiescent(Filament)
Abell 3444	428 ± 41	35.70(8.99)	66 ± 10	394 ± 140	0.67 ± 0.54	Offset
Abell 3574	161 ± 15	1.53(5.32)	0.88 ± 0.061	82 ± 16	0.56 ± 0.37	Disturbed
Abell 3581	464 ± 19	5.43(12.44)	2.4 ± 0.14	379 ± 69	1.6 ± 1.0	Disturbed
Abell 3605	269 ± 20	7.37(5.96)	4.5 ± 0.72	484 ± 130	2.4 ± 1.0	Quiescent
Abell 3638	111 ± 19	8.61(5.91)	5.9 ± 0.90	403 ± 67	1.3 ± 0.97	Quiescent
Abell 3639	150 ± 58	18.79(7.16)	56 ± 4.5	367 ± 73	1.4 ± 0.56	Quiescent
Abell 3806	136 ± 14	5.90(4.16)	3.3 ± 0.69	231 ± 52	2.9 ± 0.88	Quiescent
Abell 383	181 ± 52	24.25(7.69)	40 ± 5.0	386 ± 73	1.5 ± 0.64	Plume
Abell 3880	444 ± 20	13.24(11.72)	12 ± 0.89	438 ± 50	1.8 ± 0.95	Plume
Abell 3998	346 ± 66	17.85(10.61)	7.5 ± 0.57	253 ± 137	1.4 ± 0.18	Plume(Filament)
Abell 4059	396 ± 46	7.89(8.22)	4.1 ± 0.48	536 ± 97	2.5 ± 2.0	Quiescent

Table 2. Here we list the derived parameters for every object from the sample. The velocity column refers to the peak to peak velocity measured across the mean line of sight velocity map. The extent is the length of the major axis. The luminosity is calculated only for flux that appears at greater than 7σ significance. The FWHM is measured from the spectrum of the central $\sim 2 \times 2$ arcsec

gles to the plane of the sky. Ophiuchus had a physical offset of just 2.2 kpc but a dynamical offset of ~ 600 km s⁻¹ while Abell 1991 / 3444 have much larger physical offsets but dynamical offsets on the order of ~ 100 km s⁻¹. Abell 2566 has an offset of 5.8 kpc putting it between the two extremes. Comparing the redshift of the H α to that of the NaD from the BCG we find a dynamical offset of 210 ± 21 km s⁻¹ which is consistent with this postulation.

Abell 2566 differs from the other offset objects which are studied in Hamer et al. (2012) in that it shows a possible close companion in the VIMOS field of view (Figure 8). Comparing the NaD of the second galaxy to that of the BCG we find a velocity difference of 279 ± 29 km s⁻¹ confirming it is a cluster member and not a foreground galaxy. This suggests the possibility that the two are strongly interacting. A FORS1 R band acquisition image shows an extent of the continuum emission in the direction of the H α similar to clumpy complex structures seen in other systems

where the BCG and a companion are interacting. The offset objects previously studied showed no evidence of an interaction with another cluster member, so while Hamer et al. (2012) suggest an interaction as the likely cause of the offset it was not possible to confirm it unambiguously. Unfortunately Abell 2566 lacks high resolution X-ray observations which are required to study the ICM in detail. So, while it may be tempting to speculate that the offset H α emission in this system is evidence of sloshing induced by an interaction more detailed observations are required before this can be confirmed.

Figure 9 (top panel) shows the physical offset against the cluster redshift for the whole sample. It can clearly be seen here that the offsets for the majority of the sample fall below the seeing limit of the VIMOS observations. Abell 1991 and Abell 3444 again clearly stand out from the sample, and we see the group of plumed objects (along with Abell 2566) at an offset of ~ 6 kpc. The dot-dashed line in

Cluster	Velocity km s ⁻¹	Extent kpc(arcsec)	L _{Hα} 10 ⁴⁰ erg s ⁻¹	FWHM km s ⁻¹	[NII]/H α	Morphology
Abell 478	176 ± 12	17.04(10.62)	23 ± 2.3	261 ± 32	1.1 ± 0.92	Quiescent
Abell 496	252 ± 8	8.19(12.46)	3.1 ± 0.28	260 ± 46	2.1 ± 1.3	Quiescent
Abell 795	315 ± 38	11.13(4.66)	18 ± 3.7	755 ± 64	1.5 ± 0.85	Quiescent
Abell 85	217 ± 29	8.31(7.74)	1.6 ± 0.20	339 ± 54	2.3 ± 0.47	Quiescent(filament)
HCG62	168 ± 12	3.30(11.38)	0.087 ± 0.011	204 ± 101	3.0 ± 0.53	Quiescent
Hydra-A	556 ± 6	7.95(7.68)	13 ± 1.7	346 ± 58	0.86 ± 0.56	Quiescent
NGC4325	255 ± 10	6.66(12.98)	0.90 ± 0.089	236 ± 34	1.5 ± 0.59	Disturbed
NGC5044	277 ± 18	5.60(30.51)	0.54 ± 0.062	457 ± 99	2.5 ± 0.65	Disturbed
NGC533	211 ± 19	2.77(7.14)	0.19 ± 0.033	376 ± 218	3.1 ± 1.4	Quiescent
NGC5813	223 ± 24	1.48(10.61)	0.044 ± 0.0059	503 ± 81	2.1 ± 0.63	Disturbed
NGC5846	117 ± 12	1.27(10.12)	0.076 ± 0.0056	255 ± 63	3.0 ± 0.88	Quiescent
RXCJ0120.9-1351	345 ± 8	4.66(4.71)	1.5 ± 0.28	361 ± 40	2.2 ± 1.3	Quiescent(filament)
RXCJ0132.6-0804	452 ± 16	24.78(9.57)	43 ± 5.7	434 ± 66	1.5 ± 0.66	Quiescent
RXCJ0331.1-2100	125 ± 38	14.47(4.51)	54 ± 7.6	592 ± 184	1.2 ± 1.0	Quiescent
RXCJ0543.4-4430	245 ± 53	12.97(4.61)	4.7 ± 0.98	259 ± 169	1.9 ± 0.35	Quiescent
RXCJ0944.6-2633	289 ± 8	25.54(10.19)	36 ± 4.4	277 ± 132	0.58 ± 0.40	Disturbed
RXCJ1257.1-1339	468 ± 17	6.82(22.74)	1.4 ± 0.087	274 ± 116	1.1 ± 1.0	Quiescent
RXCJ1304.2-3030	343 ± 17	3.33(14.87)	1.2 ± 0.072	375 ± 41	1.5 ± 0.9	Quiescent(filament)
RXCJ1436.8-0900	199 ± 16	14.57(9.55)	14 ± 1.7	273 ± 14	1.4 ± 1.0	Quiescent
RXCJ1511.5+0145	138 ± 45	5.03(6.38)	0.29 ± 0.062	553 ± 113	3.0 ± 0.40	Quiescent
RXCJ1524.2-3154	310 ± 19	15.57(8.30)	46 ± 3.9	492 ± 47	1.3 ± 1.2	Plume
RXCJ1539.5-8335	333 ± 59	13.74(9.59)	30 ± 2.0	572 ± 123	1.1 ± 0.56	Disturbed
RXCJ1558.3-1410	275 ± 28	13.82(7.71)	22 ± 2.6	465 ± 103	1.2 ± 0.71	Quiescent
RXCJ2014.8-2430	539 ± 13	18.60(6.92)	140 ± 12	462 ± 92	0.78 ± 0.55	Quiescent
RXCJ2101.8-2802	169 ± 10	7.16(10.76)	1.0 ± 0.11	106 ± 17	0.54 ± 0.32	Offset
RXCJ2129.6+0005	157 ± 29	30.53(8.20)	32 ± 4.6	314 ± 64	0.93 ± 0.32	Quiescent
RXCJ2213.0-2753	185 ± 17	32.25(2.72)	0.92 ± 0.32	660 ± 372	1.4 ± 1.1	Compact/Quiescent
RXJ0000.1+0816	128 ± 3	3.55(4.57)	1.9 ± 0.17	388 ± 74	1.7 ± 0.57	Disturbed
RXJ0338+09	388 ± 6	8.60(12.54)	8.3 ± 0.50	327 ± 77	1.1 ± 0.96	Disturbed
RXJ0352.9+1941	244 ± 19	17.81(9.0)	62 ± 5.3	551 ± 45	0.79 ± 0.64	Quiescent
RXJ0439.0+0520	236 ± 54	26.15(7.70)	62 ± 7.1	559 ± 92		Quiescent
RXJ0747-19	354 ± 32	24.76(13.11)	63 ± 6.1	422 ± 83	1.7 ± 0.77	Disturbed
RXJ0821+07	270 ± 20	20.44(10.12)	35 ± 3.7	180 ± 49	0.84 ± 0.49	Plume
RXJ1651.1+0459	309 ± 25	17.50(6.52)	29 ± 4.0	276 ± 152	0.54 ± 0.39	Quiescent
S555	287 ± 20	11.54(13.08)	4.5 ± 0.44	578 ± 234	1.8 ± 1.1	Quiescent
S780	420 ± 17	33.04(8.87)	240 ± 28	490 ± 49	1.1 ± 0.95	Plume
S805	141 ± 22	2.75(8.91)	0.18 ± 0.018	227 ± 59	3.7 ± 0.95	Quiescent(filament)
S851	523 ± 17	3.58(18.56)	0.53 ± 0.039	425 ± 71	3.5 ± 3.1	Quiescent
Z3179	278 ± 15	14.80(5.92)	7.6 ± 1.5	482 ± 82	2.3 ± 0.97	Quiescent
Z348	226 ± 12	42.02(10.66)	390 ± 48	508 ± 71	0.57 ± 0.49	Quiescent

Table 2. continued.

Figure 9 marks the limit of the VIMOS field of view, assuming the BCG was at the centre of the pointing, objects with a separation placing them above this line would not be detected as the offset emission would fall outside the observed area of sky. The bottom panel of Figure 9 shows a similar comparison between the centre of the clusters X-ray emission and the H α emission for objects with high resolution X-ray observations. The X-ray centres are taken from the ACCEPT data base (Cavagnolo et al. 2009) with the exception of Abell 3444 for which we use the position defined in Hamer et al. (2012). Most of the objects have an offset close to the average seeing limit of VIMOS suggesting that there is no significant offset seen. However, we note that 3 objects have an offset of more than 4 kpc that are significantly greater than the average seeing. One is Abell

133 which shows an offset 5 kpc (4.5 arcsec). We note that the continuum image of this object shows a bright source 4.5 arcsec to the north of the BCG which matches the position of the X-ray centre suggesting that the true cluster centre may be obscured by a point source as was the case for Abell 3444 (Hamer et al. 2012). The other two objects are RXCJ 1539.5-8335 and RXCJ 2129.6+0005 which show large offsets of 8.4 kpc (6 arcsec) and 13.7 kpc (3.6 arcsec) respectively. RXCJ 1539.5-8335 may be erroneous as the VIMOS field of view only covers half of the BCG thus the position of the true peak of the H α emission is uncertain. RXCJ 2129.6+0005 is quite high redshift for the sample ($z = 0.234$) and we note that the seeing is above the average seeing limit suggesting the apparent offset may be a result of the seeing. However, the average seeing during the observations of this

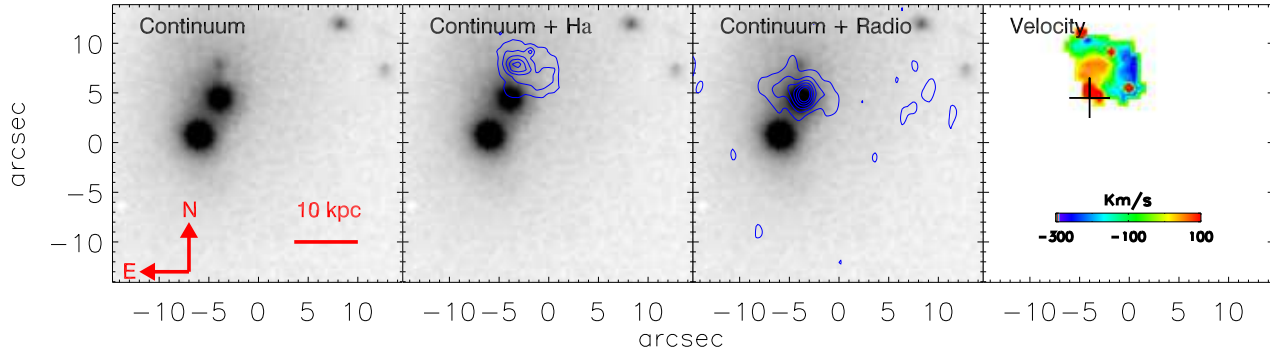


Figure 8. *Left* - FORS1 R band image of Abell 2566, the BCG is the northern of the two objects. A region of low surface brightness emission can be seen extending north from the BCG. *Centre Left* - FORS1 R band image of Abell 2566 with the $H\alpha$ emission contoured in blue. A clear offset between the BCG and the $H\alpha$ emission of ~ 6 kpc can be seen. We note that the peak of the $H\alpha$ emission is coincident with the position of the extended low surface brightness structure seen in the continuum image. *Centre Right* - FORS1 R band image of Abell 2566 with the radio map taken from Owen & Ledlow (1997) contoured in blue. *Right* - $H\alpha$ velocity field of Abell 2566 as shown in Appendix ???. The cross on this image marks the centre of the BCG.

object was 1.8 arcsec, roughly half the observed offset so it cannot be explained due to seeing alone. Additionally the position of the $H\alpha$ peak is well defined and no continuum source is present at the location of the X-ray centre suggesting that this offset may be real. Further dedicated study of these 3 objects is required to confirm if these offsets are real, however even if they all are they account for just 12% (3/24) of the objects studied. This would suggest that the optical line emission follows the X-ray peak in the majority of systems agreeing with the findings of Hamer et al. (2012).

Figures 7+9 highlight the difficulty in detecting offset emission with VIMOS observations. Due to the limited field of view large offsets at low redshift would not be detected. Likewise a smaller physical offset at the higher redshift end of our sample would have an apparent offset on the scale of the seeing (the offset in Abell 1991 would be < 3 arcsec at the redshift of Abell 3444, for example) preventing its detection. We also note that objects with an offset along the line of sight would not be easily detected when considering only the structure of the ionised gas. The discovery of other offset objects does not significantly alter the statistics discovered in Hamer et al. (2012) with just 4 of the 73 objects in this sample showing offset emission these objects account for less than 5% of the line emitting BCG population. If the sloshing of the ICM is responsible for these offsets then it suggests that events capable of causing sloshing (such as a strong interaction) are rare or do not cause a major disruption of the cold gas reservoir.

4.2 Mean velocity structure

The velocity structure of the gas can also indicate if the gas is disturbed and may be a more reliable measure in systems where the lowest surface brightness gas may not have been detected. Maps of the mean line of sight velocity (from here on the mean velocity) structure of each object were produced from the parameterised fits so the velocity structure could be studied. The ionised gas in many of the cluster cores studied show a clear velocity gradient from negative to positive velocities relative to the systemic velocity. This velocity gradient is often seen to run across the bright peak

of the line emission. Velocity profiles like these are often indicative of rotation but can also arise from uniformly entrained or out flowing material with an inclination to the plane of the sky. One caveat that must be considered here is the possibility that limited resolution in the higher redshift objects could smooth out variations in the velocity field and make the overall variation across the BCG appear as a much smoother gradient than is actually the case. To test this we take NGC 5044, a low redshift ($z \sim 0.009$) object with a velocity field that shows a lot of variation, and degrade its resolution to that of a $z=0.05$ object before refitting the cube and remaking the velocity maps. While this significantly reduces the small scale variations in the velocity field seen in the higher resolution map the overall velocity structure still does not show a smooth gradient like those seen in other objects suggesting that limited resolution has only a small impact on our ability to distinguish between objects with a smooth velocity gradient and those without.

We define the peak-to-peak velocity to be the maximum velocity difference measured across the natural peak of the line emission. To calculate this the velocity maps were first median smoothed to eliminate any pixels with extreme velocities relative to their surrounding pixels. A ~ 2 arcsec wide slit was then placed on the map centred on the peak of the line emission, in the case of objects with significantly offset emission we use the peak of the continuum instead. This slit was then rotated through 180° in 6° increments and at each position a profile was extracted by interpolating along the slit and smoothing to 2 arcsec to reduce the noise. Figure 10 shows an example profile from Hydra-A with the blueshift peak, redshift peak and peak-to-peak velocity indicated. The maximum velocity difference between the blueshift peak and the redshift peak (relative to the redshift at the centre of the $H\alpha$ emission) from these profiles is then defined as the peak-to-peak velocity. This is the maximum velocity difference across the line emission as given in Table 2. One caveat to note is that the blueshift and redshift peaks were required to be on opposite sides of the centre. In some of the more disturbed velocity maps (NGC 5044 for example) the natural peaks did not occur on opposite sides of the centre, in such cases the peaks were defined as the

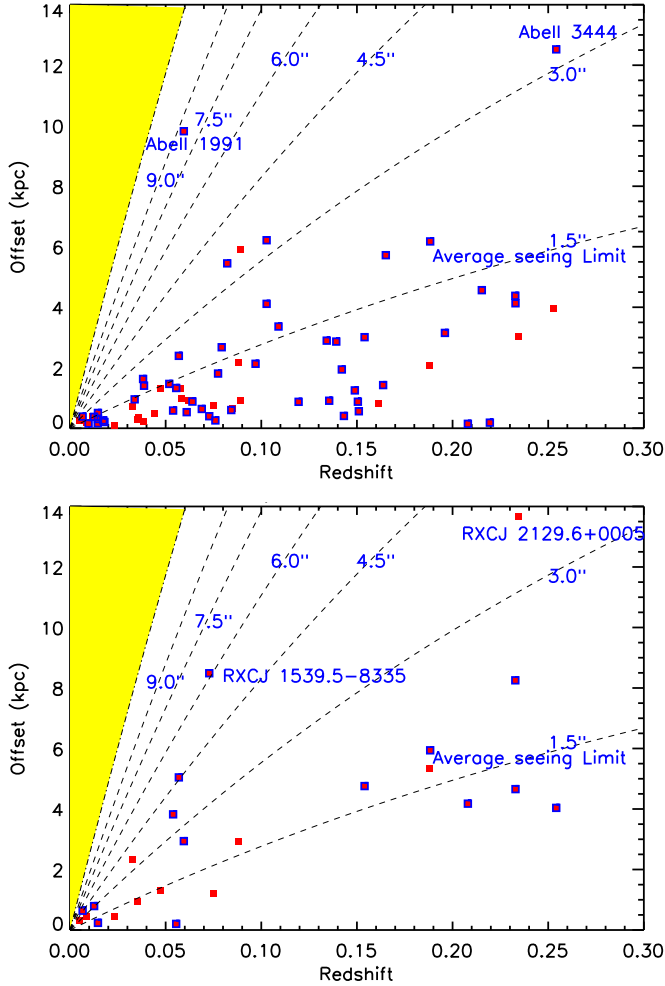


Figure 9. *Top* - Observed physical offset between the BCG and the majority of the line emission at the cluster redshift. *Bottom* - Observed Physical offset between the centre of the clusters X-ray emission and the majority of the line emission at the clusters redshift. The dashed lines show constant visible offsets in $1.5''$ steps. We can see that the majority of objects fall well below the average seeing limit of $1.5''$. The dot-dashed line shows the extent of the VIMOS field of view ($\sim 12''$ assuming the BCG was positioned at the centre of the field of view) with the shaded region marking the extent beyond the field of view. Objects whose seeing during the VIMOS observations was less than $1.5''$ (the average seeing limit) have their points highlighted by a blue outline.

points which gave the maximum velocity difference across the centre of the $H\alpha$ emission.

Within the more extended and disturbed objects the velocity maps show much more variation. However, under closer scrutiny many of these objects do still often show smooth velocity gradients on smaller scales close to the centre of the system (defined by the location of the $H\alpha$ peak). The bright central region of these systems show a coherent velocity structure with a strong gradient across the centre of the system. Away from this central region however, in the lower surface brightness extended emission the velocity field becomes much more chaotic. Abell 3581 and RXJ 0338+09 are two good examples of such objects, each show extended low surface brightness emission with no coherence to the ve-

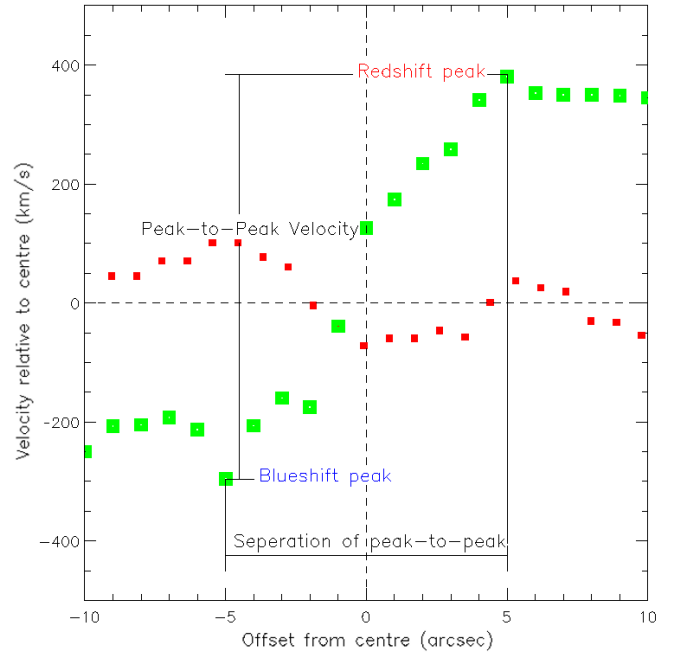


Figure 10. Two examples of a velocity profiles taken from Hydra-A (large green points) and NGC 5044 (small red points). These profiles were used to calculate the peak-to-peak velocity within each system. The contrast of these two objects shows how the peak-to-peak was clear in some objects (such as Hydra-A) while being much more difficult to constrain in others (such as NGC 5044). The blueshift peak and redshift peak are labelled and the peak-to-peak velocity and peak-to-peak separation are indicated for Hydra-A.

locity structure. However, in the regions close (within ~ 8 arcsec) to the peak of the emission the velocity field shows much less variation. If it is the case that the extended emission is the result of some event which has disturbed the distribution of the cold gas then this small scale velocity structure may represent the overall velocity field of the object. Events such as entrainment by radio jets or galaxy interactions which can disturb the cold gas would initially affect the low density gas at the edges of the distribution. The disturbance can impart a velocity on this gas which is not consistent with the overall velocity field of the system. The densest gas however, is much more difficult to displace (Churazov et al. 2002) and, if it remains relatively undisturbed, should retain its previous velocity, which should be governed by the global velocity field within the cluster core.

If the velocity field of each object is roughly consistent, then any observed peak-to-peak velocity difference between them would simply be the effect of the gas sampling different regions of the velocity distribution. This would present itself as a clear trend between the extent of the line emission in a given object and the magnitude of the velocity difference observed.

In Figure 11 we show the peak-to-peak velocity plotted against the maximum continuous extent of the line emission. A lower velocity difference is seen in very compact objects, however, as very compact objects are much more susceptible to blurring caused by the seeing this trend may not be real. If the velocity difference is being measured over a re-

gion comparable to the seeing it is possible that gas at each peak may be blended with some emission from the other peak resulting in a shifting of the best fitting models closer together thus reducing the measured velocity difference. For objects with an extent of > 2 kpc the median velocity difference of the objects remains constant at ~ 250 km s $^{-1}$ with a very large scatter. It is important to note however, that any inclination of the velocity structure towards the plane of the sky will reduce the observed peak-to-peak velocity.

The average velocity gradient of each system can be calculated from the peak-to-peak velocity difference and the separation of the peak-to-peak (thus $V_{grad} = \Delta V / \Delta r$).

In Figure 12 we plot the calculated values of velocity gradient for each object, the top two panels of this figure show the velocity gradient plotted against the two parameters used to calculate it (the peak-to-peak velocity and the separation of the velocity peaks). It can be seen from these plots that the measured velocity gradient of each object is well correlated with the separation with a smaller gradient in objects with a larger Δr . By contrast the plot showing the velocity gradient against the velocity difference shows no obvious correlation. It is not unexpected that a correlation would be seen when plotting $\Delta V / \Delta r$ against Δr . The simple expectation if ΔV and Δr are uncorrelated would be that $V_{grad} \propto \Delta r^{-1}$ and $V_{grad} \propto \Delta V$. However, our best fit trend line has the equation $V_{grad} = 160 \pm 1.2 \times \Delta r^{-0.8 \pm 0.06}$ which is not consistent with this expectation. No significant trend is found in the plot of V_{grad} against ΔV , however we note that higher values of the peak-to-peak velocity have (on average) higher velocity gradients. This is shown by the green diamonds which show the median velocity gradient within 50 km s $^{-1}$ bins. It can be seen that at higher peak-to-peak velocities the median velocity gradient is higher, as would be expected. The red dashed line shows the expected trend $V_{grad} \propto \Delta V$ with the form $V_{grad} = \Delta V / 10$, the points of median velocity follow this trend (except at extreme velocities where few points are sampled) suggesting that the

velocity gradient is proportional to the peak-to-peak velocity as is expected, albeit with considerable scatter.

In order to test the expected correlations we replace the peak-to-peak velocity (ΔV) in our analysis with the total H α flux of an object. Since the H α flux comes from the full extent of each object we expect it to have no correlation to the separation of the peak-to-peak velocity. We thus define $H\alpha_{grad} = F_{H\alpha} / \Delta r$ and plot $H\alpha_{grad}$ against $F_{H\alpha}$ and Δr . We find that $H\alpha_{grad} = 62 \pm 1.5 \times \Delta r^{-0.94 \pm 0.15}$ and $H\alpha_{grad} = 0.11 \pm 1.4 \times F_{H\alpha}^{0.97 \pm 0.09}$. Since the exponents of these equations are consistent with -1 and 1 respectively the equations are consistent with our expectation that the total H α flux has no dependence of the separation of the peak-to-peak velocity. Thus for two unrelated values a and b, $a/b \propto a$ and $a/b \propto b^{-1}$. Since the correlation of velocity gradient ($\Delta V / \Delta r$) is proportional to $\Delta r^{-0.8}$ not Δr^{-1} (Figure 12) this suggests that ΔV already has some dependence on Δr that is a physical phenomenon, rather than a mathematical artefact. As such the velocity fields have some structure as a function of radius.

One possible explanation is that the velocity field traces a rotation. On large scales gas rotating under gravity should form a Keplerian disc about the centre of mass of the system. The rotation curves of discs are well understood and can be approximated to follow the curve of an arctan function (Courteau 1997) of the form

$$V(r) = V_0 + \frac{2}{\pi} V_c \arctan(R) \quad (2)$$

where $R = (r - r_0)/r_t$, V_0 is the velocity at the centre of rotation, r_0 is the spatial centre of the galaxy, V_c is the asymptotic velocity corrected for inclination ($V_c = V_{asym} \times \sin i$) and r_t is the radius at which the rotation curve transitions from rising to flat (where the rising part of the rotation curve has a gradient of 1). Thus for a rotating disc the velocity at a given distance from the centre of rotation can be assumed to be a function of that distance.

If the velocity field followed such a rotation then the expected peak-to-peak velocity (ΔV) would be $2 \times$ the rotational velocity ($V(r)$ from Equation 2) and the separation (Δr) would be twice the maximum radius (r from Equation 2). V_0 can be set to 0 as we are only interested in the velocity difference not the absolute velocity. We used Equation 2 to produce 500 velocity profiles using randomly generated variables from within the parameter limits of our data (V_{asym} of 200–700 km s $^{-1}$, $\sin(i)$ of 0.0–1 and r_t of 1–3 kpc). For each we calculated the velocity gradient (V_{grad}) and reproduced Figure 12 for these models. We find a correlation between V_{grad} and the peak-to-peak velocity with the form $V_{grad} = 130 \pm 1.1 \times \Delta r^{-0.78 \pm 0.03}$ and no strong correlation of V_{grad} with ΔV which matches what is seen in the data. This suggests that the velocity structure in the majority of our objects may be consistent with disc like rotation. However, we note that the velocity maps for some of our objects are very chaotic (e.g. NGC 5044) and clearly not dominated by rotation, thus a more detailed analysis is needed to identify rotation in these systems on a case by case basis. We will further explore the idea that the kinematics are related to a rotating disc in the next paper in this series but have already shown this to be true for one object (Hamer et al. 2014).

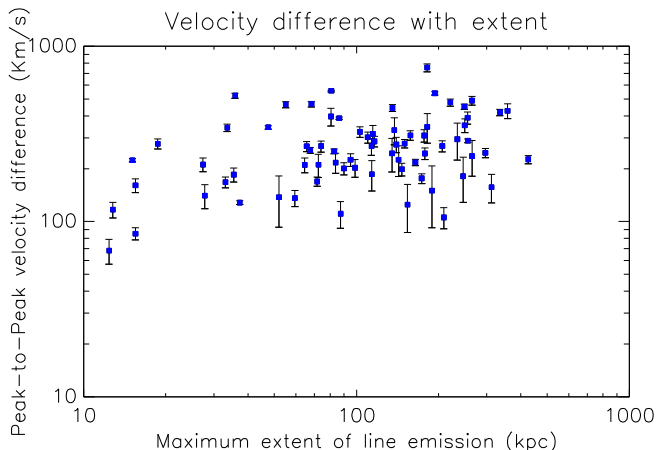


Figure 11. The peak-to-peak velocity difference against the measured extent of each object. It is clear that there is no trend with the exception that very compact objects show a reduced velocity difference. This suggests that the different peak-to-peak velocities seen between objects in the sample are not simply the effect of having gas sample a larger portion of the velocity field in some objects than in others.

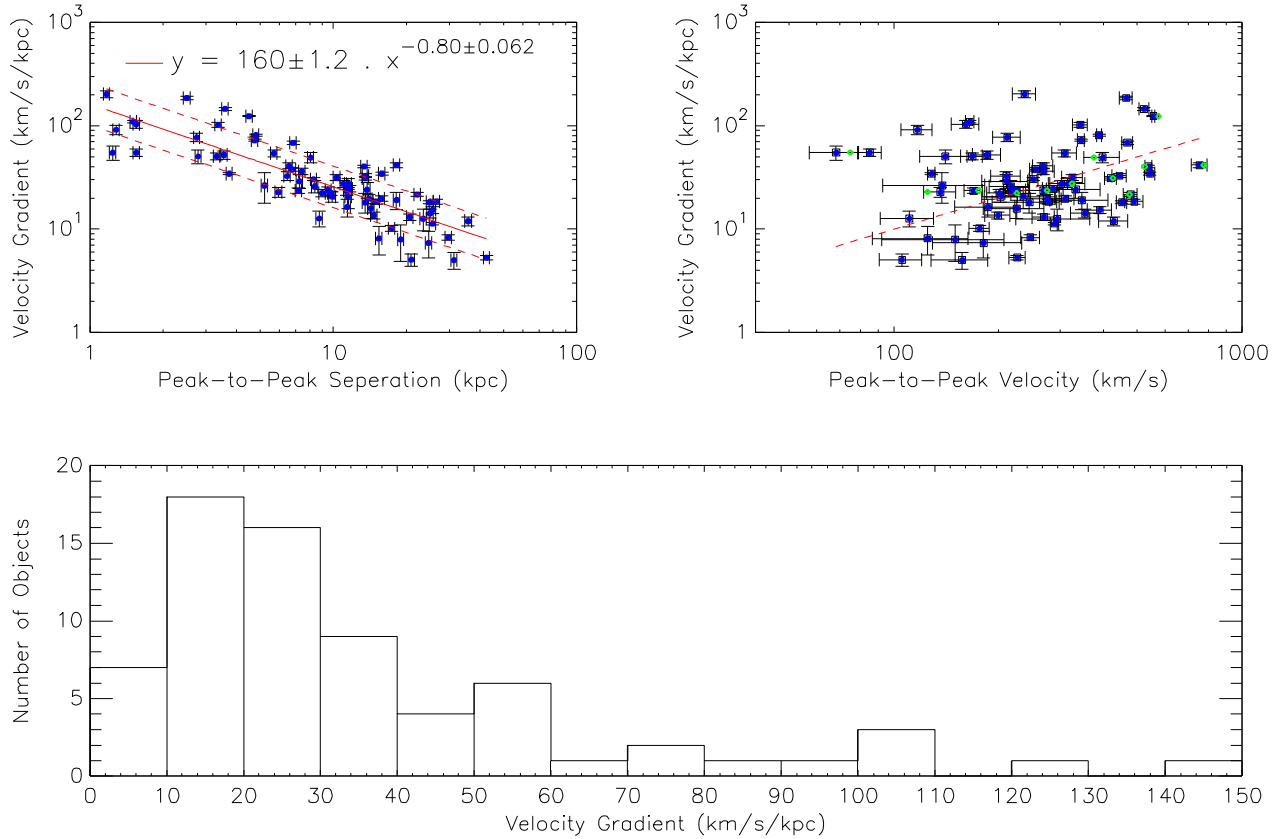


Figure 12. *Top Left* - Plot of the velocity gradient of every object in the VIMOS sample against the separation of the peak-to-peak velocity. A clear trend can be seen in this plot indicated by the red line and the equation of the best fit is given. *Top Right* - In this plot we compare the velocity gradient of our objects to the measured peak-to-peak velocity difference. No clear trend is seen, the dashed line shows a linear increase with velocity and the green diamonds show the moving average of the data. *Bottom* - This shows a histogram of the velocity gradients calculated for the objects in the VIMOS sample. A clear peak can be seen at $\sim 15 \text{ km s}^{-1} \text{ kpc}^{-1}$ with a steady decrease towards higher velocity gradients. This suggests that the velocity gradients seen in the sample are sufficiently small that their effect on the velocity dispersion should be negligible. The combined interpretation from the top two plots is that the velocity gradient responds to the separation of the peak-to-peak in a non linear way while showing little correlation to the peak-to-peak velocity itself.

4.2.1 Velocity dispersion

The FWHM of an emission line is a measure of the velocity dispersion of the gas from which it was emitted. In observations with ordered motions (such as a difference in the mean velocity) the velocity dispersion is a measure of the random motions of the particles which make up the gas. There is of course some care required in this interpretation as it requires the assumption that the variation of any ordered velocity field, within the region the spectrum is extracted from, is negligible compared to the random velocities within that same region. Figure 12 also shows a histogram of the mean velocity gradients for each object in the sample. From this it can be seen that the velocity gradient for most objects is quite low with the distribution peaking at $\sim 15 \text{ km s}^{-1} \text{ kpc}^{-1}$. This suggests that the velocity gradients seen in the sample are sufficiently small that their effect on the measured linewidths should be negligible.

We give the average line width measured within the central $\sim 2 \times 2 \text{ arcsec}$ region of the system in Table 2. We note that the gradient of the clusters velocity field should be highest in the centre where the mass concentration naturally peaks and it is within this region that the velocity dispersion

is most likely to be broadened by the velocity gradient. The region used to extract the FWHM is comparable to the mean seeing ($\sim 1.5 \text{ arcsec}$) and is thus the scale on which we would expect the mean velocities to be blended.

In Figure 13 we show the average velocity dispersion within the central $2 \times 2 \text{ arcsec}$ of each objects plotted against the peak-to-peak of the mean velocity field. If we ignore the points with a peak-to-peak velocity below 100 km s^{-1} (i.e. below the spectral resolution of the **HR_Orange** and **HR_Red** VIMOS gratings) this plot would suggest that there is no obvious correlation between the average central velocity dispersion and the peak-to-peak velocity difference for our sample. We performed a Spearman's test to confirm this lack of correlation and find that $\rho = -0.016$ with a p-value of 0.9 confirming the lack of correlation. This lack of correlation confirms the assumption that the velocity dispersion measured is real and not an effect of blending gas clouds with different mean velocities. Higher spectral and spatial resolution data would be required to study the broad line regions in more detail.

The maps of most objects show a FWHM profile which peaks towards the centre of the system, where the flux map peaks, and falls away as the surface brightness of the line

emission does. This line width profile is exactly what is expected as in the central regions there should be a higher projected mass density. Within most objects the extended low surface brightness gas has a consistently low FWHM on the order of $\sim 100\text{--}200\text{ km s}^{-1}$. Interestingly some objects show additional peaks in the line width away from the centre suggesting the gas is more kinematically disturbed in these regions. Indeed the fact that most of these offset FWHM peaks appear in systems with spatially disturbed morphologies (Abell 3574, RXJ 0338+09) would seem to support this assumption.

4.3 Statistics of the sample

In Figure 14 we plot the $\text{H}\alpha$ luminosity against the extent for each object. As would be expected we see a clear trend in this plot with the more luminous objects being more extended due to the increased surface brightness. The points on the plots in Figure 14 are colour coded, in the top plot by the peak-to-peak velocity, and in the middle by the $[\text{NII}]/\text{H}\alpha$ ratio (see Section 5 for a discussion of the line ratio analysis). The top plot of Figure 14 shows no general trend of the velocity, as might be expected from the lack of a trend in Figure 11. We do note however that the lowest luminosity ($< 10^{40}\text{ ergs s}^{-1}$), most compact objects (extent $< 3\text{ kpc}$) have low peak-to-peak velocities with all but two falling into the lowest velocity bin. The middle plot of Figure 14 shows how the $[\text{NII}]/\text{H}\alpha$ ratio varies with the extent and luminosity. Again we note that the objects to the lower left (compact and low luminosity) have quite consistent $[\text{NII}]/\text{H}\alpha$ ratios with all but two falling into the highest bin. So collectively from Figure 14 we can identify a collection of objects with low extent and luminosity, little velocity structure and high ionisation states. We would therefore suggest that the majority of the objects in Figure 14 below an $\text{H}\alpha$ luminosity of $\sim 4 \times 10^{39}\text{ ergs s}^{-1}$ with an extent less than 4 kpc are

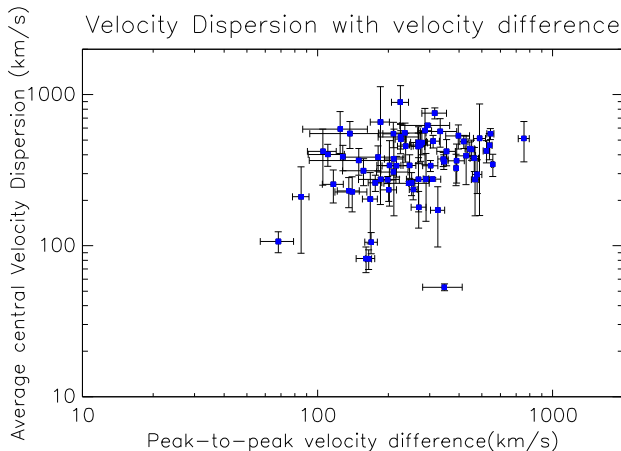


Figure 13. Plot of the central velocity dispersion (FWHM) against the peak-to-peak velocity difference of each object. If the outliers with very low velocity dispersion ($< 100\text{ km s}^{-1}$) and/or low velocity difference ($< 100\text{ km s}^{-1}$) are ignored then there is no obvious trend suggesting the velocity dispersion is unrelated to the velocity across the object. This confirms the assumption that the velocity dispersion measured is real and not an effect of blending gas clouds with different mean velocities.

most likely to be LINERS. The bottom plot of Figure 14 colour codes the points by redshift bin. The higher redshift sources are typically brighter and more extended as would be expected for a flux limited sample.

In Figure 15 we compare each objects $\text{H}\alpha$ luminosity with its peak-to-peak velocity. There is no trend immediately apparent from this plot as might be expected given the trend between luminosity and extent, and apparent lack of trend between velocity and extent. When considering an observed velocity structure it is always important to consider the effects of inclination which can reduce the observed peak-to-peak velocity regardless of the origin of the velocity structure. The most likely inclination to view a velocity structure along, given a heterogeneous population and unbiased sample (a fair assumption for our objects given the minimal selection criteria), is 60 degrees to the plane of the sky. If we assume that the average velocity for each luminosity bin is from an object at this inclination, then we can show how this trend would be affected by the inclination. We selected a number of inclinations which would split the velocity range approximately into quarters, these inclinations are plotted as a series of lines in Figure 15. We can see that the same trend shifted to 90° encloses all of the objects in the sample bar one, S851. If we ignore the single outlier then it is possible to account for the different velocities seen in objects with similar luminosities as purely an effect of inclination.

If we assume this distribution in velocities to be a result of the inclination of the system then we can calculate the apparent inclination angle of each system using its luminosity and peak-to-peak velocity. In Figure 16 we show a histogram of this apparent inclination for the systems in the sample. This plot shows that the velocity structure of the systems are distributed about an apparent inclination of 60° to the plane of the sky. There are very few objects with an apparent inclination around 0° and few with an inclination of greater than 80° . This distribution is consistent with what is typically seen when calculating the inclination of galaxies from their axial ratios. It is important to note however, that using such a method has degeneracy with the thickness and ellipticity of the galaxy (Tempel et al. 2013). Thus edge on and face on disc galaxies have their inclination incorrectly calculated resulting in a distribution similar to that in Figure 16. However, the method used to determine the inclination in Figure 16 is not subject to these same degeneracies and thus the distribution should be constant with $\sin(\theta)$. The lack of face on systems can be explained by the fact that the measurable true peak-to-peak velocity difference would be very small. In such a situation the velocity field will be dominated by noise obscuring face on systems. We also note that some scatter in velocity at a given luminosity is expected, especially given the size of the luminosity bins used (an order of magnitude). So the highest velocity systems in a given bin may skew the mean thus making naturally lower velocity systems appear more inclined resulting in an under-abundance of edge on systems.

4.4 Comparison to the Stellar Kinematics

The sodium D stellar absorption features ($\text{NaD}_{\lambda 5895.9}$ and $\text{NaD}_{\lambda 5889.9}$) allows for the extraction of stellar kinematics but it requires the continuum emission to be bright in order

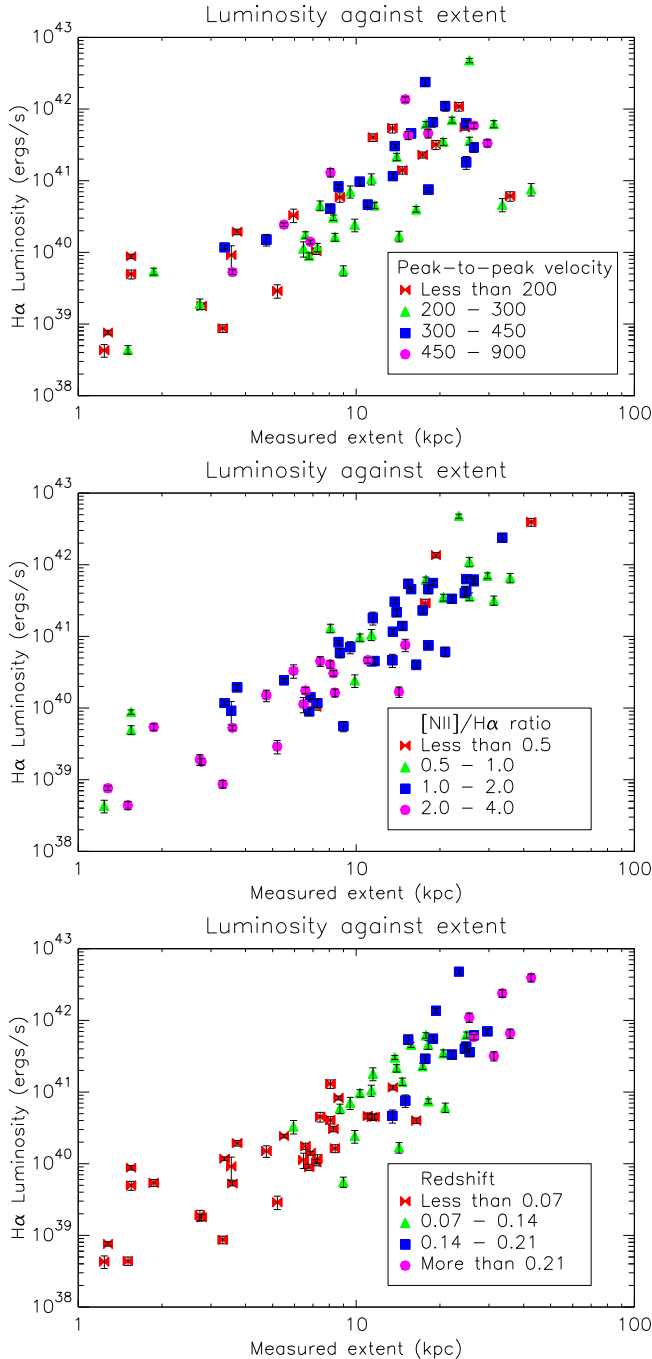


Figure 14. The $H\alpha$ luminosity plotted against the extent of the line emitting region. A clear trend can be seen between the luminosity and extent as would be expected. *Top* - The points are colour coded into bins of peak-to-peak velocity. The velocity shows no clear trend on this plot however, we do note that the most compact (with an extent of less than 3 kpc) and least luminous objects have low velocities with all but 2 falling into the lowest velocity bin. *Middle* - The points are colour coded into bins of overall $[NII]/H\alpha$ ratio. The trend of $[NII]/H\alpha$ ratio with extent can be seen in this plot. Again we see that the low luminosity compact objects typically have high $[NII]/H\alpha$ ratios. *Bottom* - The points are colour coded into redshift bins. The higher redshift sources typically show higher luminosities as would be expected for a flux limited sample.

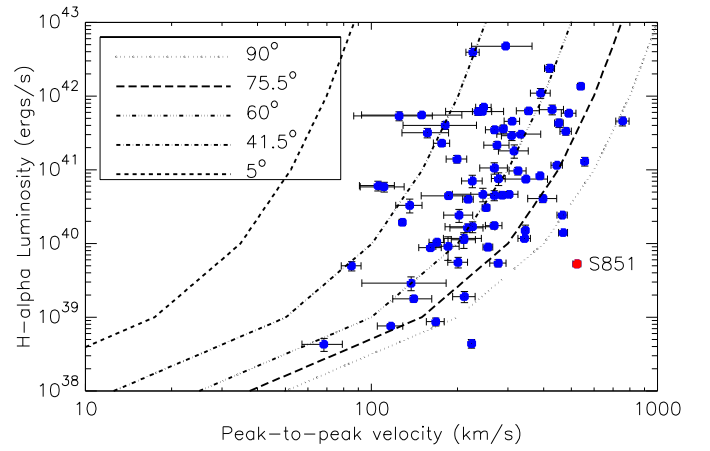


Figure 15. The $H\alpha$ luminosity plotted against the peak-to-peak velocity. It is apparent from this plot that the mean velocity of the sample is dependent on the luminosity. This may be an effect of the more $H\alpha$ luminous systems being more massive or that we can detect $H\alpha$ sampling more of the velocity field in the more luminous systems. The lines on this plot show what the velocity would be at a given inclination if we assumed this trend was true, took the median velocity in each luminosity bin and assumed this came from an object inclined at 60 degrees to the plane of the sky (the most likely inclination of an unbiased sample). In such a case the velocity of only one object in the sample, S851, cannot be explained by an inclination effect alone.

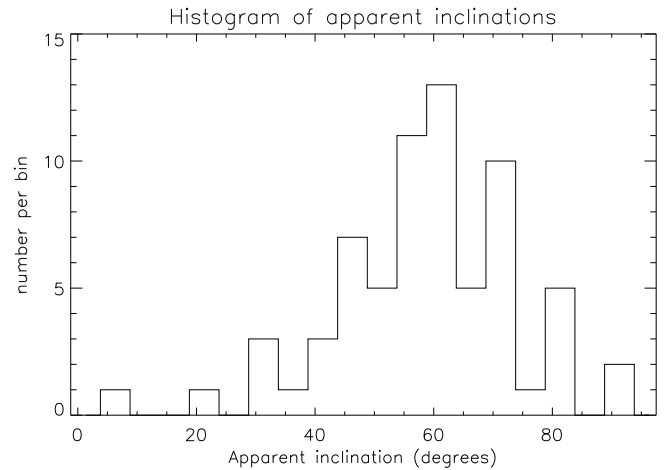


Figure 16. Here we show the apparent inclination of the objects in the sample if it is assumed that the variation in their velocity at a given luminosity is a function only of inclination. The apparent inclination of the velocity is relative to the plane of the sky such that a velocity in the plane of the sky would have an inclination of 0 degrees. There is a peak at ~ 60 degrees around which the apparent inclinations are distributed likely caused by noise obscuring face on systems and degeneracy between the effect of inclination and natural variations of the peak-to-peak velocity.

for the absorption feature to be visible in the spectrum. It is also intrinsically a broad feature which makes contamination by sky lines a serious problem. We performed the fitting using Voigt absorption profiles following the same procedure used to fit the Gaussian emission line profiles. From the sample only 12 objects had spectra in which the NaD absorption

was clear enough to be fitted at the 7σ significance level (so as to match the emission lines).

While these objects all show a strong continuum and strong NaD absorption, they are otherwise quite dissimilar in their X-ray and optical properties. They have X-ray luminosities in the range of $0.1 - 42.10 \times 10^{43} \text{ erg s}^{-1}$ and $\text{H}\alpha$ luminosities in the range of $0.087 - 46 \times 10^{40} \text{ erg s}^{-1}$ which is representative of most of the objects in our sample. We note that none of the most luminous systems (either in X-ray or $\text{H}\alpha$) appear in this subsample. However, such objects make up only a small fraction of our sample. The 12 objects show no common features in their other optical properties (velocity, extent, FWHM, $[\text{NII}]/\text{H}\alpha$ ratio and morphology) and were observed in a wide range of seeing conditions ($0.63 - 2.39 \text{ arcsec}$) suggesting they are representative of the sample. We do note that the objects are all typically low redshift sources (11 out of 12 have $z < 0.06$) but this is the result of the requirement for bright continua to detect the NaD absorption feature.

Recent studies have separated the NaD absorption in star forming galaxies into two components (one from the stars and one from the ISM) by fitting stellar population models to isolate the stellar component (Heckman et al. 2000; Jiménez-Vicente et al. 2007). This may be possible for the brighter objects in our sample through binning to increase the signal to noise. However, we expect the stellar component to be dominant in our systems as NaD is strong in the spectrum of cool stars (Jacoby et al. 1984) common in BCGs and is one of the strongest optical stellar absorption lines in the spectra of early type galaxies (Heckman 1980; Bica et al. 1991). As such, given the focus of this paper is on the ionised gas emission we assume that all the NaD absorption is stellar in origin.

We produce maps from the fits to the NaD absorption feature for these 12 objects and compare them to the line emission maps (see Appendix ??). Determining the velocity dispersion of the NaD absorption from these line fits is not a trivial calculation and the fits themselves constantly overestimated the FWHM. To account for this we used the fitting method to fit the total spectrum of each object and compared the FWHM derived by this method with the FWHM derived by fitting stellar templates to the same spectrum using penalised pixel-fitting (Cappellari & Emsellem 2004, which could not be used on individual pixels due to insufficient signal to noise). In this way we calculated a normalisation factor for the FWHM in each object and applied this to the velocity dispersion maps.

One thing that is immediately apparent from studying the NaD absorption maps is that the stellar component of the BCGs shows no evidence of an ordered velocity structure. Indeed the mean velocity field appears to be completely random which is expected for a massive elliptical galaxy such as a BCG. This is in stark contrast to the ordered velocity maps produced from the $\text{H}\alpha$ emitting gas. We also note that although the line width maps do generally show a peak towards the centre of the object they are overall substantially more uniform than their counterparts from the line emission. This implies that the stellar component of the BCGs has a velocity field which is dominated primarily by random motion suggesting that the movement of the stars and gas in the galaxy are decoupled. In Figure 17 we compare the kinematics of the gas to those of the stars. No trends can

be seen for any of the comparisons in this figure however we note the large errors on the parameters from the stellar fits make this ambiguous.

To test this we performed a Monte Carlo procedure to reproduce the plots with parameters sampled from within the Gaussian errors shown. We produced 10,000 of each of the four comparisons: stellar velocity against gas velocity, stellar velocity against gas FWHM, stellar FWHM against gas velocity and stellar FWHM against gas FWHM. For each we performed a Spearman correlation test and found no correlation in 90.0%, 87.6%, 88.9% and 89.2% of the realisations respectively. This lack of correlation confirms that the kinematics of the ionised gas are decoupled from those of the BCGs stellar component. Hydra-A is a clear example of this, it is one of the most ordered velocity fields in our sample and also has one of the largest peak-to-peak velocities. Moreover it is known that the gas forms a rotating disc (Simkin 1979; Hamer et al. 2014), despite this however the stellar motions appear to remain random, producing a velocity field with no obvious order.

In Section 4.1.1 we identify Abell 1991 as an object with offset line emission. The extent of this offset can be seen in the maps comparing the stellar absorption to the ionised gas emission which can be found in Appendix ???. While both have a peak in roughly the same position, the line emission map shows little emission around this point with most of it located $\sim 10 \text{ arcsec}$ to the north where it shows a second peak. By contrast the majority of the NaD absorption is centred almost uniformly around the southern

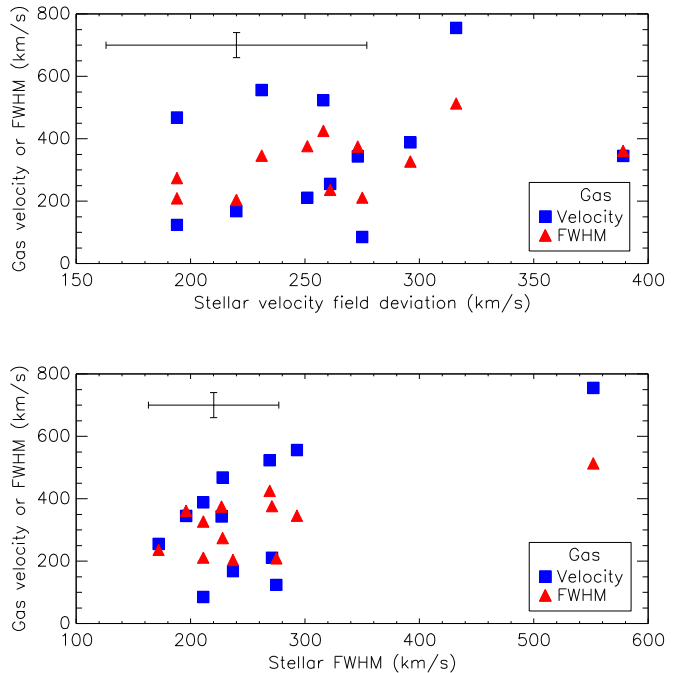


Figure 17. *Top* - Ionised gas peak-to-peak velocity and FWHM compared to the standard deviation of the stellar velocity field. *Bottom* - Ionised gas peak-to-peak velocity and FWHM compared to the overall FWHM of the stellar absorption. No trend can be seen in either of these plots suggesting that the ionised gas kinematics are decoupled from the stellar component of the BCG. Representative error bars are shown in the top left of each plot.

peak of the line emission with possibly a slight extent to the south. The comparison of the line emission maps with the stellar absorption maps clearly highlights this offset

The line emission map of RXJ0338+09 shows a very disturbed morphology suggesting that some major event is disrupting the line emitting gas within this system. Interestingly the stellar absorption map shows a secondary peak of absorption at ~ 7 arcsec (~ 4.7 kpc) to the north-west of the BCG. This suggests the presence of a second galaxy very close in projection to the BCG. Consulting the stellar velocity map we find that this galaxy also has a velocity which is close to that of the BCG, redshifted by just a few hundred km s^{-1} and matching the velocity of the line emission at this point. It is tempting to interpret these findings as evidence of an interaction which is disrupting the cold gas in the cluster core. However, we note that the companion galaxy has a bright X-ray point source (Sarazin et al. 1992) suggesting AGN activity so it is possible that the $\text{H}\alpha$ emission seen at the redshift of the companion is associated to its nucleus and superimposed on that of the cluster core. Additionally the ICM shows evidence of cavities from multiple AGN outbursts (Sanders et al. 2009) which may be responsible for disrupting the gas on large scales. Finally the rotationally excited molecular gas appears to form a rotating structure close to the BCG (Wilman et al. 2011). Thus to fully characterise the nature of the “interaction” in this system requires a dedicated study of the complex processes at work using high sensitivity multiwavelength observations and is beyond the scope of this paper.

5 THE PHYSICAL PROPERTIES OF THE GAS

In order to investigate the spectral differences between the objects we now present spectra extracted from two regions for each object in our sample. We initially extracted spectra from the full region of line emission for each object. We defined this extent as being the region where the $\text{H}\alpha$ flux exceeded the noise by a factor of 7. To determine this we first collapsed the cube (see Section 2 for a description of this procedure) over the region of the spectra containing the $\text{H}\alpha$ emission. Additionally, we collapsed the cube over two other regions of equal spectral length which contained no emission lines or sky lines to produce two continuum images. A continuum subtracted $\text{H}\alpha$ map was then produced by subtracting the first of these continuum maps from the $\text{H}\alpha$ map. The first continuum image was also subtracted from the second continuum image to produce a map from which the noise could be determined.

A mask was then produced to exclude spatial pixels where the $\text{H}\alpha$ emission did not exceed a Signal to Noise (S/N) of 7. The flux in the remaining pixels was then summed for each channel to produce a spectra of the full $\text{H}\alpha$ emitting region. A second spectra for each object was extracted from the nuclear region of the BCG, defined as a 3×3 pixel ($\sim 2 \times 2$ arcsec²) region (roughly the extent of the mean seeing) centred on the peak of the continuum emission. All spectra were sky and continuum subtracted, and have been filtered to eliminate contamination by cosmic rays. The two spectra from each object can be found in Appendix ?? and the line fluxes and ratios are given in Appendix ??.

Our IFU observations of these clusters allow us to study the physical properties of the gas within the ionised nebulae using the ratios of the various optical emission lines visible in these spectra. Many of the lines commonly used in such analysis ($\text{H}\alpha$, [NII], [OI] and [SII]) are covered by our observations for all objects and are detected in most (see Table 1). Importantly as our observations are resolved they allow us to determine how the properties of the gas vary throughout the nebula. However, some care must be taken with this analysis, especially as we lack the [OIII]/ $\text{H}\beta$ ratio (commonly used in conjunction with the other line ratios for this kind of analysis), as there are many factors which can all affect the observed line ratios.

5.1 Metallicity

The line ratios are subject to change due to variations in the metallicity of the gas with the [NII]/ $\text{H}\alpha$ being the most strongly affected, $\log_{10}([\text{NII}]/\text{H}\alpha)$ varies linearly with metallicity until it saturates at $\log_{10}([\text{NII}]/\text{H}\alpha) \sim -0.5$. For cool core clusters metallicity profiles of the ICM are consistent between objects (Böhringer & Werner 2009), falling from close to solar metallicity in the BCG to $\sim 1/3$ solar metallicity (Loewenstein 2004) in the ICM (beyond $\sim 0.25 r_{200}$). We thus expect that the metallicity of the gas is approximately the same in all of the objects and, as our observations only cover at most a few 10s of kpc, that any change in metallicity within a given object will be small. Thus the metallicity should have a negligible effect on the variation of line ratios both between objects and within a single object when compared to other effects.

As all objects in our sample have $\text{H}\alpha$ and [NII] emission we can test this assumption by using the $N2$ index to estimate the metallicity of the gas in each of our objects as measured by the oxygen abundance (O/H). We use equation 1 from Pettini & Pagel (2004) to estimate the range of oxygen abundances in our sources from the [NII]/ $\text{H}\alpha$ ratio. From Table 2 we see that the maximum and minimum mean [NII]/ $\text{H}\alpha$ ratios in our sample are 3.7 and 0.54 respectively which corresponds to oxygen abundances ($12 + \log_{10}(\text{O}/\text{H})$) of 9.22 and 8.75. Pettini & Pagel (2004) find that 95% of their measured oxygen abundances lie within ± 0.41 of the value derived for a given [NII]/ $\text{H}\alpha$ ratio. Our two extremes fall well within these limits suggesting that the metallicities within all our objects are consistent. For objects where we also have the [OIII] and $\text{H}\beta$ lines we can further test this consistency by using the $O3N2$ index as an additional measure of the oxygen abundance. Using equation 3 from Pettini & Pagel (2004) we estimate the oxygen abundances for the seven sources for which we could obtain the [OIII]/ $\text{H}\beta$ ratios. We find oxygen abundances ($12 + \log_{10}(\text{O}/\text{H})$) of 8.67 to 9.16 which are consistent with the values derived from the [NII]/ $\text{H}\alpha$ ratios. As such we assume that the metallicity is the same in all objects and has no effect on the measured line ratios.

5.2 Electron Density

The density of the gas can also have an effect on the line ratios measured for a given excitation mechanism. Importantly models of star formation excitation often used to distinguish between star forming galaxies and those excited by

more energetic phenomena typically assume electron densities of a few hundred or less (for example Dopita et al. 2000; Kewley et al. 2001). The strengths of several lines in our observations are sensitive to the electron density but in particular the [SII] doublet line fluxes, available for all but a few objects in our sample, are sensitive to the effects of collisional de-excitation. Since both lines have very similar excitation energies their relative excitation rates are almost insensitive to variations electron temperature. As such the effect of temperature is negligible and the relative excitation rates of the two lines depends only on the the ratio of their collision strengths and thus the ratio of intensities depends only on the density of the gas (Osterbrock 1989). This analysis is only sensitive to densities over ~ 1 – 2 orders of magnitude between 100 cm^{-3} and 10^4 cm^{-3} , fortunately this is consistent with the densities assumed in models of star formation excitation and we expect the electron density in the cores of clusters to be towards the lower end of this density range. In Figure 18 we plot the $[\text{SII}]_{\lambda 6716}/[\text{SII}]_{\lambda 6731}$ ratio against the $[\text{SII}]_{\lambda 6731}$ flux for all objects in the sample. The dashed horizontal lines on this plot contain the region that is sensitive to density changes. The majority of the points are found towards the upper boundary of this region suggesting that most objects in our sample have electron densities on the order of a few 100 cm^{-3} . We also note however that a small number of points (just four) fall within the region below a value of ~ 0.7 suggesting these objects have densities in excess of 10^3 cm^{-3} .

The fact that some points fall outside of the region sensitive to density must also be considered. Some of these can be accounted for by the fact that in some cases some or all of the $[\text{SII}]_{\lambda 6716}$ line fell into an atmospheric absorption feature and was lost. An example is Abell 795 which has a ratio in its core of ~ 0.2 despite relatively strong lines being present as the $[\text{SII}]_{\lambda 6716}$ line is redshifted into the atmospheric A-band. A similar situation also affects some of the objects which have a ratio above 1.45. In this case the $[\text{SII}]_{\lambda 6731}$ line falls into the absorption feature causing the $[\text{SII}]_{\lambda 6716}$ to appear unusually strong. Some of these points can also be attributed to objects whose redshifts put the [SII] lines right at the edge of the spectral coverage of our observations. In some cases this resulted in some of the $[\text{SII}]_{\lambda 6731}$ line falling off the edge of the spectrum in part or all of the field of view. Objects with such obvious factors causing errors in the [SII] flux measurements have been removed from the analysis in Figure 18. There remain however, many objects with ratios falling above a threshold of 1.45 and still two that fall below a ratio of 0.45. We note however, that the errors on the ratio are considerable, in Figure 18 we show representative errors positioned at the median ratio of $[\text{SII}]_{\lambda 6716}/[\text{SII}]_{\lambda 6731}$. We note that the upper error limit encompasses most of the points above a ratio of 1.45, as the majority of our objects tend to have high $[\text{SII}]_{\lambda 6716}/[\text{SII}]_{\lambda 6731}$ ratios (above ~ 0.8) these errors can explain the large number of points falling above our expected upper limit. Indeed the majority of ratios fall within these errors suggesting that they are consistent with the median ratio of 1.32 which corresponds to densities on the order of 100 – 200 cm^{-3} . As such the measured gas densities are consistent with those used in models of star formation so we can use the expected line ratios from such models to determine if star formation plays a major role in the excitation of the ionised gas in cluster cores.

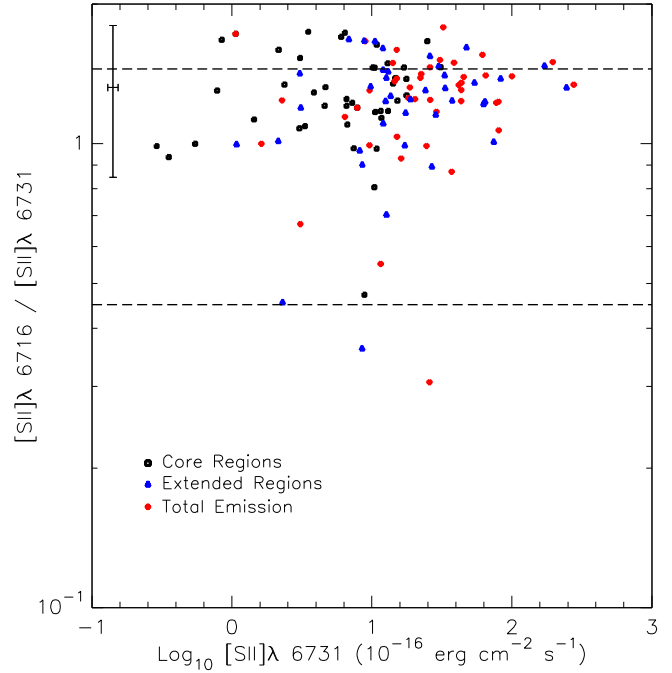


Figure 18. A comparison of the $[\text{SII}]_{\lambda 6716}/[\text{SII}]_{\lambda 6731}$ ratio and $[\text{SII}]_{\lambda 6731}$ flux for all objects in the sample which had detections of both lines. The points show the core regions (black squares), extended regions (blue triangles) and total emission (red diamonds) for each object. The dashed lines indicate upper and lower bounds where the $[\text{SII}]_{\lambda 6716}/[\text{SII}]_{\lambda 6731}$ ratio is sensitive to changes in electron density. The upper boundary is set to a value of 1.45 corresponding to a density of $\sim 100 \text{ cm}^{-3}$ and the lower boundary to a value of 0.45 corresponding to a density of $\sim 10^4 \text{ cm}^{-3}$. Most of the points fall towards the upper boundary suggesting densities commonly in the region of a few 100 cm^{-3} .

5.3 Excitation state

The [NII] to $\text{H}\alpha$ ratio is also a good measure of the excitation state of the gas and in particular is a good indicator of even small contributions to the excitation from an AGN thanks to its sensitivity to metallicity. The $\log_{10}([\text{NII}]/\text{H}\alpha)$ ratio saturates at high metallicity peaking at ~ -0.5 (Denicoló et al. 2002; Kewley & Dopita 2002; Pettini & Pagel 2004; Kewley et al. 2006), as such [NII]/ $\text{H}\alpha$ ratios of above ~ 0.6 are a strong indicator of an AGN (or other non star formation) contribution to the excitation. Table 2 lists the mean [NII]/ $\text{H}\alpha$ ratio for all objects in our sample as calculated from the fits to each pixel. The ratio varies substantially between objects suggesting that some objects are substantially more highly ionised than others. As noted earlier some of the objects showed $\text{H}\alpha$ emission which was absorbed, this will have the effect of artificially rising the [NII]/ $\text{H}\alpha$ ratio calculated from fits to the total spectrum. We clearly indicate which objects show $\text{H}\alpha$ absorption in Table 1. For these objects the [NII]/ $\text{H}\alpha$ ratio calculated and presented in Table 2 can at best be considered an upper limit.

For the majority of objects the [NII]/ $\text{H}\alpha$ ratio can be seen to vary within the maps presented in Figure 4 suggesting that the ionisation state is not constant throughout the objects. This is not surprising as sources of ionisation (AGN, star formation, shocks, etc.) are often seen to be localised

to specific parts of a galaxy (AGN are localised to the centre of the galaxy, star formation is often seen to be clumpy rather than uniformly distributed O’Dea et al. 2010). The maps also indicate the regions of objects which have substantial $H\alpha$ absorption as the $[NII]/H\alpha$ ratios are extremely high appearing white on the maps.

While the maps show a lot of variation in the ionisation state within different regions of each object, the line ratios do not show any obvious structures. Studying the overall ionisation state allows for a direct comparison between objects and with other global parameters. The ionising radiation within galaxies is believed to come from two primary sources, AGN and star formation (Baldwin et al. 1981). AGN typically produce a harder radiation field which results in a higher ionisation state while star formation has a softer radiation field. Thus one would expect gas ionised by AGN to have a high $[NII]/H\alpha$ ratio while that ionised by star formation should have a lower value. However, this distinction is more complicated for gas in BCGs where both AGN and star formation are likely to be occurring. Shocks, produced from either internal (e.g. AGN outbursts) or external (e.g. galaxy interactions) sources, can also produce optical line emission. Generally, faster shocks produce higher line ratios (though there is also a dependence on metallicity, gas density and magnetic fields, Allen et al. 2008) comparable to those produced by AGN. However, if such fast shocks were present they should be visible in Figure 4 as broad regions in the FWHM panel. Post-AGB stars have been shown to be a possible source for LINER like emission (Singh et al. 2013) which can produce line ratios inconsistent with star formation models. Other possible sources of the ionisation include photoionisation by emission from the hot gas in the ICM or collisional heating (Ferland et al. 2009; Fabian et al. 2011) which can also produce high $[NII]/H\alpha$ ratios.

The majority of the extracted spectra show little variation in appearance between the central region and the total emission suggesting that an AGN is not significantly affecting the distribution of ionising radiation in most objects. However, some objects do show a variation in the $[NII]$ to $H\alpha$ line ratio, which might suggest the influence of an AGN. Indeed for some objects the $[NII]$ to $H\alpha$ ratio is sufficiently high (above $[NII]/H\alpha \sim 0.6$) to rule out star formation as the dominant component of the ionisation. One striking observation from the spectra is the fact that some objects, which show strong lines in their total spectrum, have little or no emission present in their central regions. These are objects in which the peak of the line emitting gas is offset from the BCG which results in weak or absent lines at the peak of the continuum emission. Such objects are discussed in section 4.1.1 and studied in detail in Hamer et al. (2012) and while interesting, make up only a small fraction of BCGs (just 4 of the objects in our sample are offset in this way) and are thus not representative of the population.

In Figure 19 we show the global measure of the $[NII]/H\alpha$ ratio of each system as a function of the extent of its line emitting region. With the exception of a few outliers this plot shows a trend with the $[NII]/H\alpha$ ratio decreasing as the extent increases. Since the ionising radiation in BCGs likely to be generated from numerous sources (AGN, star formation, soft X-rays from the ICM, cosmic rays, shocks, collisional excitation etc.) this trend can be interpreted as an effect of the relative contributions of each. Radio obser-

vations of all BCGs in this sample have already been taken with a substantial number detected. Hogan et al. (2015) find a flat spectrum radio source at 10 GHz in 36 of the 73 objects in our sample, and a steep spectrum component at 1 GHz in 53/73 of our objects. This suggests that AGN activity is currently active, or has recently been active within the cores of many of the BCGs in our sample. As such in the very compact objects the gas is likely to be concentrated near an AGN so the ionisation will thus be dominated by fast shocks and photoionisation producing a high $[NII]/H\alpha$ ratio. The more extended objects will have a higher relative contribution from other less energetic sources of ionisation producing a lower overall ratio explaining the global trend we see in Figure 19. In the next paper in this series we will use resolved radio observations to compare the radio structure to the kinematics of the ionised gas to determine how the gas and AGN interact.

5.4 The source of the ionisation

We have already discussed the ionisation state of the gas in the previous section, however it is important to note that while the $[NII]/H\alpha$ ratio can provide a measure of how highly ionised any gas is it cannot on its own be used to determine the source of the ionising radiation. For this, other line ratios are also required, such as $[OIII]/H\beta$ for a BPT analysis (Baldwin et al. 1981). Unfortunately for most of the objects in our sample the observations did not cover the wavelength range required to observe these lines. However, as can be seen from Table 1, only seven of the objects in our sample were at a redshift high enough to allow $H\beta$ and $[OIII]$ to be present at the blue end of the HR_orange grism.

We produced BPT diagrams for the 7 objects (Figure 20) which had both $H\beta$ and $[OIII]$ line present in the spectrum. From Figure 20 it is apparent that all 7 of these ob-

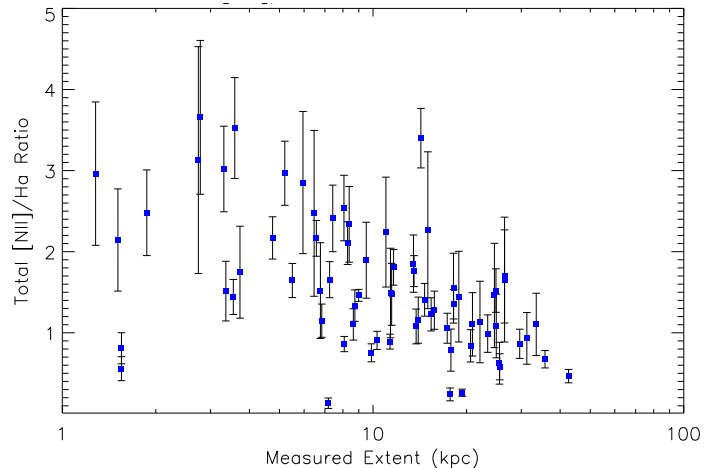


Figure 19. The overall $[NII]/H\alpha$ ratio for every object in the VIMOS sample plotted against the extent of the line emitting region. A loose trend can be seen in this plot which suggests that the overall $[NII]/H\alpha$ ratio decreases as the extent of the emission line region increases. This suggests that the dominant source of the ionisation is more energetic for compact objects than for extended ones. Likely the compact objects feel more influence from a central AGN while the extended regions are ionised by other mechanisms.

objects appear to have ionisation states which are not consistent with being the result of star formation activity based on the empirical separation from Baldwin et al. (1981) (dotted line). It is interesting to note that one of the objects for which the BPT analysis was possible was Abell 2566. Abell 2566 was discussed in Section 4.1.1 because the emission is completely offset from the BCG by ~ 9 kpc. Kewley et al. (2001) calculate the maximum theoretical boundary for star formation which is shown as a dashed line in Figure 20. We note that Abell 2566 falls below this line, within the region enclosed by the two dividing lines which Kewley et al. (2006) define as a region of composites between AGN and SF ionisation. Most of the points sampled in the other 6 of our 7 objects fall above or on this line suggesting that star formation alone cannot explain the ionisation. These objects show line ratios typically attributed to AGN. While this is not unexpected as BCGs do often host AGN (Russell et al. 2013; Hlavacek-Larrondo et al. 2013) there are also other processes which could produce the line ratios seen, shocks (Allen et al. 2008) or particle heating (Ferland et al. 2009; Fabian et al. 2011) for example. Additionally we note that the $\log_{10}([\text{OIII}]/\text{H}\beta)$ values typically fall below ~ 0.5 putting these objects in the region typically occupied by LINER-like objects which Singh et al. (2013) suggest do not require an AGN as the ionising source. If AGN are responsible for ionising the gas in all these systems then this would suggest that the influence of the AGN is able to reach far beyond the core of the galaxy.

While the $[\text{OIII}]$ and $\text{H}\beta$ lines were not covered by all of our observations, the $[\text{OI}]$ and $[\text{SII}]$ lines appear in many of the observations by virtue of being close in wavelength to the principal $\text{H}\alpha$ line. While these lines are not as good an indicator of the mean level of ionisation and temperature as $[\text{OIII}]/\text{H}\beta$, their ratio with $\text{H}\alpha$ are indicators of the relative importance of high-energy ionisation in large partially ionised regions (Osterbrock 1989). These lines were detected in the extracted spectrum in the majority of our sample (see Appendix ?? for a list of line fluxes) allowing us to make some measure of the importance of high-energy ionisation sources within the cores of clusters. In Figure 21 we show histograms of the ratio of $[\text{NII}]$, $[\text{SII}]$ and $[\text{OI}]$ with $\text{H}\alpha$ and note that most ratios fall above the division between low and high energy ionisation (Osterbrock 1989; Kewley et al. 2006).

In Figure 22 we plot the $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ ratio against the $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$ ratio. The $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$ ratio shows a clear separation between sources ionised by low energy sources (such as UV emission from star formation) and those ionised by other sources (such as AGN or shocks). Kewley et al. (2006) shows this divide occurs at approximately $\log_{10}([\text{OI}]_{\lambda 6300}/\text{H}\alpha) = -0.9$ with all their sources above this value being from non stellar processes. This is shown as a vertical line on Figure 22. Similarly for the $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ ratio, Osterbrock (1989) shows a divide at approximately $\log_{10}([\text{NII}]_{\lambda 6583}/\text{H}\alpha) = -0.2$. While this separation is less obvious since at higher values of $\log_{10}([\text{OIII}]_{\lambda 5007}/\text{H}\beta)$ lower values of $\log_{10}([\text{NII}]_{\lambda 6583}/\text{H}\alpha)$ are seen it is still true that all objects above this value in the Osterbrock (1989) sample are from AGN. Thus without access to $[\text{OIII}]$ or $\text{H}\beta$ we use it as a first order estimate of the separation and show it in Figure 22 as a horizontal dashed line.

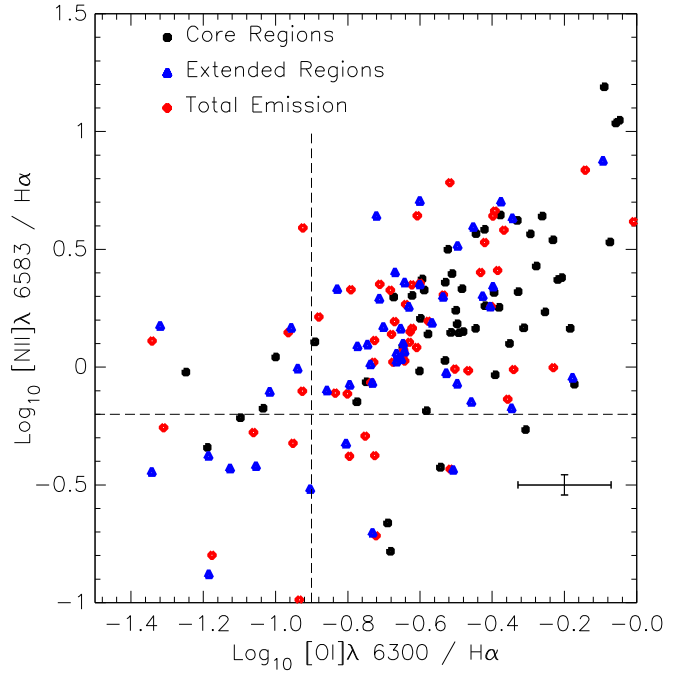


Figure 22. A comparison of the $[\text{NII}]/\text{H}\alpha$ and $[\text{OI}]/\text{H}\alpha$ ratios for all objects in the sample which had detections of all three lines. The points show the core regions (black squares), extended regions (blue triangles) and total emission (red diamonds) for each object. The dashed lines indicate the level above which empirical evidence suggests ionisation is caused by non stellar processes such as AGN or shocks. The majority of sources fall within the upper right hand quadrant, this is especially true for the core regions which show only 9 objects falling below one of these lines and only 2 falling beneath both.

It is immediately apparent from Figure 22 that the majority of the points fall in the upper right quadrant suggesting they are high ionisation regions dominated by excitation from non star formation processes. It is not surprising then that a lower fraction of the core regions fall below these cut offs than the regions containing extended or total emission (9/73 compared to 13/73 and 14/73 respectively).

We show a comparison of the $[\text{SII}]_{\lambda 6716+6731}/\text{H}\alpha$ ratio against the $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$ ratio in Figure 23. Like the $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$ ratio, $[\text{SII}]_{\lambda 6716+6731}/\text{H}\alpha$ shows a clear separation between sources ionised by high energy photons (such as from narrow line AGN) and those ionised by other sources (Osterbrock 1989). Like the $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ the separation point shows some variation at high values of $[\text{OIII}]_{\lambda 5007}/\text{H}\beta$ but as the BPT diagrams in Figure 20 suggest this is generally not the case in our sample we adopt the value $\log_{10}([\text{SII}]_{\lambda 6716+6731}/\text{H}\alpha) = -0.4$ as the separation between high-energy ionisation regions and low-energy ionisation regions. This is the maximum value that a HII region had in the sample of Osterbrock (1989), above this value all sources were ionised by non stellar processes. We plot this separation as a horizontal dashed line on Figure 23 along with the same separation for the $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$ ratio as used earlier. Once again the objects show a clear trend towards higher values of both ratios suggesting the ionisation is dominated by high energy processes such as AGN or shocks. Like for $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ we see a lower fraction of core regions (7/73)

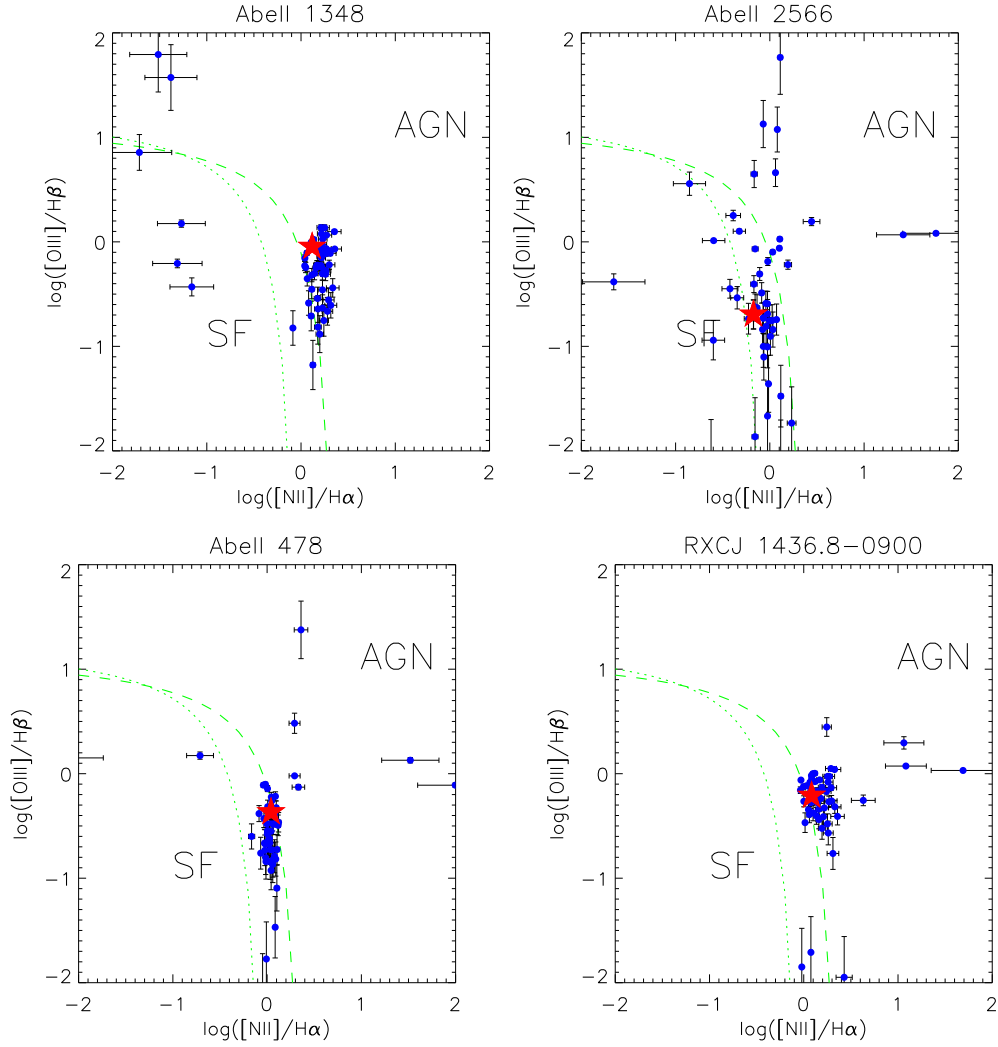


Figure 20. Here we show the BPT analysis of the objects from our sample which have redshifts that allow us access to the [OIII] and H β lines in our VIMOS observations. The blue points show the locations of every pixel of detected emission on the diagram, the red stars then show the location of the central 2×2 arcsec² which should be the location of the AGN. The green dotted line marks the empirical boundary between star formation and AGN dominated ionisation while the green dashed line marks the theoretical upper limit of star formation dominated ionisation. The region between these two lines likely contains objects which are a composite of star formation and AGN ionisation. Note that for Abell 2566 the emission was offset so the central region was determined from the H α emission and not the continuum.

fall below these cut offs than the regions containing extended (15/73) or total (15/73) emission.

In Figure 24 we compare the $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ and $[\text{SII}]_{\lambda 6716+6731}/\text{H}\alpha$ ratios. Once again the dashed lines represent the separation between high-energy ionisation and low energy ionisation regions. The majority of the points fall above both cut offs again (with only 8/73 cores, 9/73 extended regions and 12/73 totals falling below the cut offs) suggesting that the gas is excited by more energetic processes. Interestingly the points above both separating values show a much tighter correlation than is seen when either ratio is compared with $[\text{OI}]_{\lambda 6300}/\text{H}\alpha$. The reason for this is not clear however, we do note that $[\text{OI}]_{\lambda 6300}$ is the weakest line used for this analysis and as such has larger measurement errors than the other lines. This may be responsible for introducing additional scatter in the other comparison plots but we note that there is considerably more scatter below

the cut off values in Figure 24, which is comparable to that seen in the other comparison plots.

Using the results from all three of these plots we can categorise the 73 objects in our sample (separated into total emission, core emission and extended emission) based on all three line ratios. We classify star forming objects as those whose ratios fall below the cut off values for all three ratios, AGN like objects (those excited by AGN or other high ratio producing mechanisms) have all three ratios above the cut off values, ambiguous objects have at least one ratio above and at least one below the cut offs and unclassified objects have one line missing (due to absorption, falling outside the spectral coverage or just being too weak to detect) preventing analysis using all three ratios. According to this scheme our sample contains 5 objects whose total emission is consistent with star formation, separating the cores and extended emission shows that only 3 cores are dominated by star for-

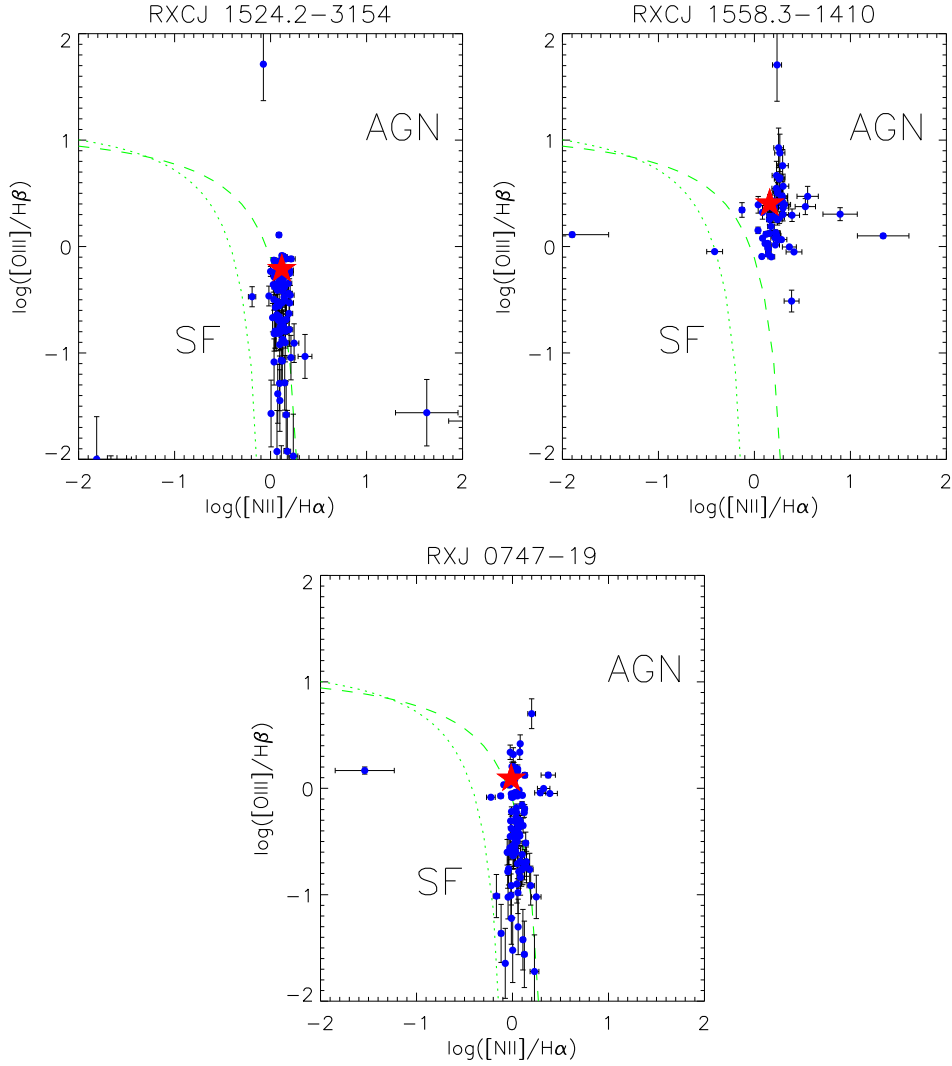


Figure 20. continued.

mation while the number of extended regions rises to just 6. By contrast, 46 objects have a total emission consistent with AGN like excitation, with 47 cores and 41 extended regions also showing AGN like excitation. We find that 18 objects have ambiguous ratios in their total emission, with 10 cores and 16 extended regions also being ambiguous when analysed separately. Finally we find the emission of 4 total regions, 13 cores and 10 extended regions cannot be classified due to having at least one line missing from their spectrum.

Some care must be taken when interpreting these ratios, in particular the presence of any extreme outliers. Due to the range in redshift some objects in the sample showed atmospheric absorption features in the region of their spectra containing these principle lines. For some objects this was very apparent as a dip of the spectra to negative values. However, this was not always the case especially in systems with strong lines. Fortunately, the relatively narrow absorption feature and large redshift range mean that this problem only affects a handful of objects in the sample. We do not include objects affected by this absorption in our plots thus allowing us to draw general conclusions from the overall statistics.

Importantly, the majority of the points fall above the transition values between low and high-energy ionisation regions for all lines suggesting that the gas ionisation is not dominated by star formation within the BCG. This is consistent with our BPT analysis presented earlier and with previous studies of line ratios in cluster cores (Johnstone & Fabian 1988; Crawford & Fabian 1992).

AGN are expected to be present in the cores of BCGs (Russell et al. 2013; Hlavacek-Larrondo et al. 2013) and are thus a potential source of the ionisation seen, especially in the core regions. However, it is important to note that the extended regions of the ionised nebula in clusters can reach out over 10s of kiloparsec from the core of the BCG. This would require the AGN to remain the dominant ionising radiation source over many orders of magnitude in scale. Shocks could be present through out the gas and can produce the line ratios seen. However, comparing to the shock models of Allen et al. (2008) our objects would require shocks over a larger range of velocities ($100\text{--}1000\text{ km s}^{-1}$) to fully explain all the line ratios seen. We would expect high velocity shocks to be visible as regions of enhanced

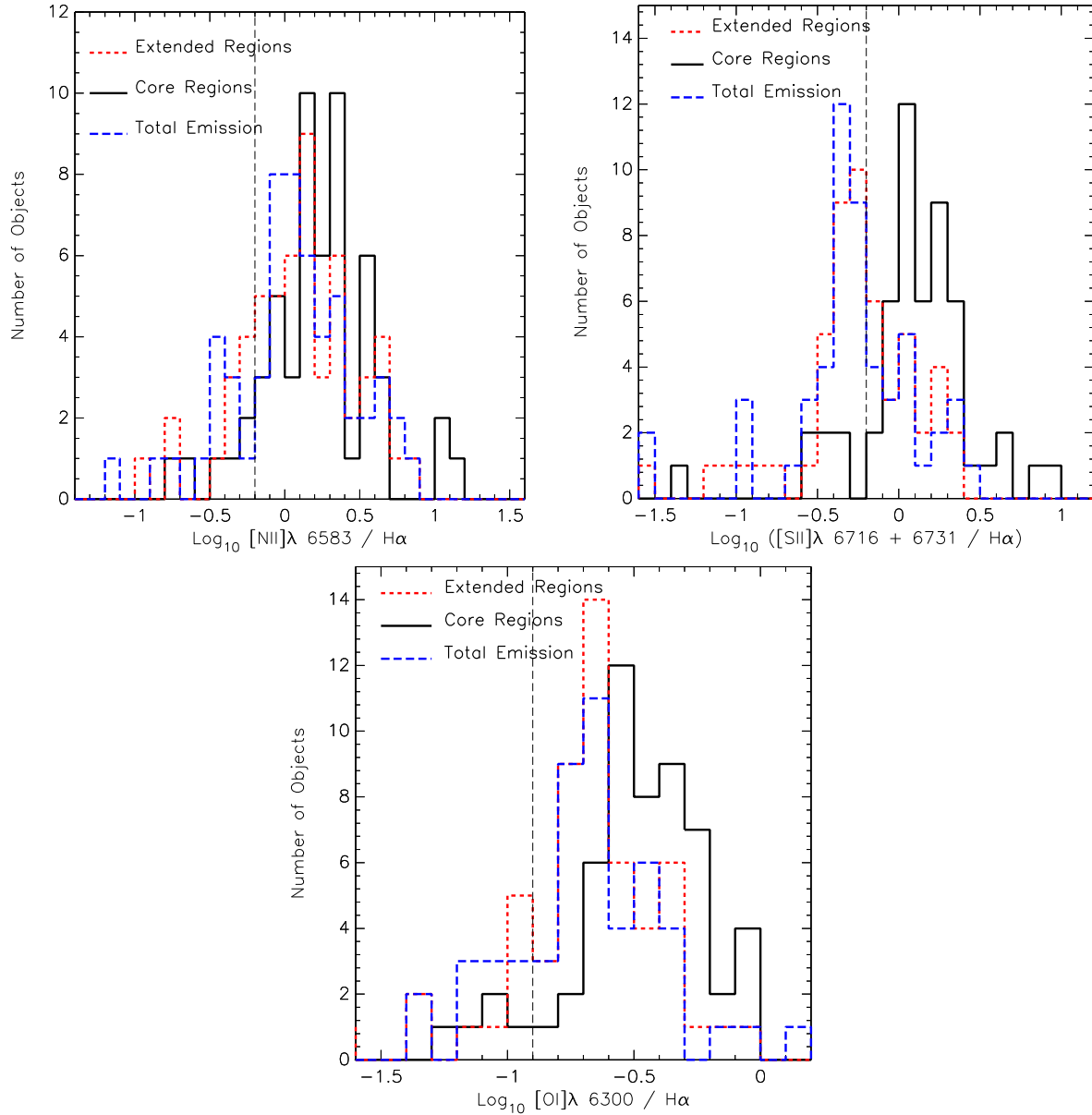


Figure 21. Histograms of the [NII] (left), [SII] (right) and [OI] (bottom) to $H\alpha$ ratios for all object in the sample. The ratios from the total emission as well as from the extended and core regions are shown separately to allow for easy comparison. The vertical long dashed line shows the division between low and high energy ionisation for each line ratio. Note that core regions typically fall at higher ratios suggesting they are ionised by more energetic sources such as AGN or shocks implying an AGN is at least partially responsible for the elevated ratios.

line width. While the central regions of our objects do show line widths (FWHM) sufficiently broad to be consistent with such shocks most objects have FWHM on the order of ~ 100 – 200 km s^{-1} in their extended regions suggesting that the fast shocks needed to produce the line ratios seen are not present here. The line ratios also do not agree well with the shock models of Allen et al. (2008). Of the seven objects with [OIII] and $H\beta$ detections only one has an [OIII]/ $H\beta$ ratio high enough to be consistent with a shock in excess of 300 km s^{-1} . From Figure 21 we see that the majority of the $\log_{10}([NII]/H\alpha)$ ratios fall within 0.0 – 0.5 which are consistent with those expected for shocks (~ 0.0 – 0.4) however, the $\log_{10}([OI]/H\alpha)$ ratios are mostly below -0.3 making them in-

consistent with the values expected for shocks (~ -0.3 – 0.2). The $\log_{10}([SII]/H\alpha)$ ratios for the majority of extended regions (-0.5 – 0.0) are consistent with those expected for shocks (~ -0.5 – 0.0) but the core regions are not (0.0 – 0.4). As such there is no consistent evidence that shocks play a major role in the excitation of the gas.

Other sources of excitation within cluster cores (particle heating for example, Ferland et al. 2009; Fabian et al. 2011) are therefore required to contribute significantly to the excitation of the gas. This can explain how the line ratios can remain relatively consistent throughout the gas, over a range of different scales, linewidths and average line ratios. This also explains how the ionised gas in Abell 2566 shows

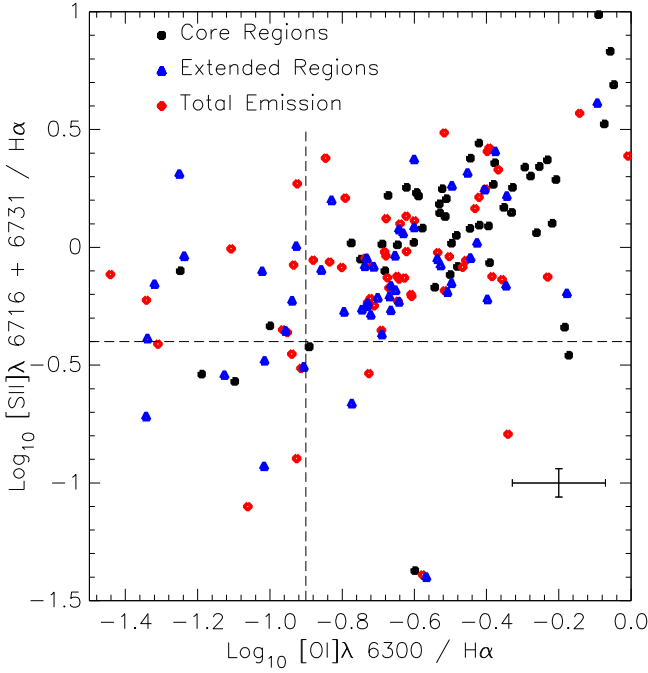


Figure 23. A comparison of the $[\text{SII}]/\text{H}\alpha$ and $[\text{OI}]/\text{H}\alpha$ ratios for all objects in the sample which had detections of all three lines. The points show the core regions (black squares), extended regions (blue triangles) and total emission (red diamonds) for each object. The dashed lines indicate the level above which empirical evidence suggests ionisation is caused by non star formation processes. The majority of sources fall within the upper right hand quadrant, this is especially true for the core regions which show only 7 objects falling below this for one of these ratios and only 2 falling beneath both.

evidence of AGN like ionisation despite being offset from the BCG (and thus, far from the AGN) and having narrow line widths. While AGN, shocks and star formation may not be the main source of the ionisation they are likely present in at least some of our objects leading to local variations of line ratios within a given object and can thus explain why some objects do show local variations while others do not. In particular the trend towards higher average line ratios in more compact objects seen in Figure 19 suggest that while other mechanisms are likely dominating the excitation of the gas globally, the AGN still plays a role.

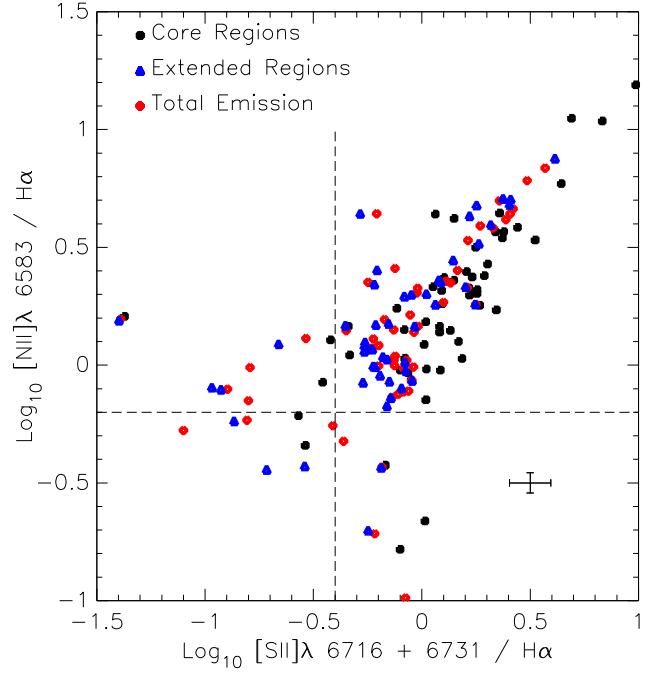


Figure 24. A comparison of the $[\text{SII}]_{\lambda 6716+6731}/\text{H}\alpha$ and $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ ratios for all objects in the sample which had detections of all four lines. The points show the core regions (black squares), extended regions (blue triangles) and total emission (red diamonds) for each object. The dashed lines indicate the level above which empirical evidence suggests ionisation is caused by non star formation processes. The majority of sources fall within the upper right hand quadrant, suggesting that non stellar processes dominate the excitation.

6 DISCUSSION

6.1 What fraction of line emitting BCGs are highly disturbed

Our analysis showed that the extent of the line emitting gas in cluster cores varies widely from system to system. A few objects are barely resolved by VIMOS, their extent dominated by the seeing of the observations with their line emission concentrated at the centre of the BCG. From the full sample of 73 objects only Abell 194 appeared to be truly compact (extent \leq seeing) in all directions. A further two were compact along a single axis but showed extent along the other. However, the majority of the objects had a true extent on the order of 5 to 20 kpc with a few extending out to 50 kpc (S780, Z348 etc.) to where the line emission far outshines the stellar continuum.

From the full sample we identify five distinct morphological states. The compact objects described above account for the smallest fraction. While it is possible that these objects do have structure, it would require higher resolution observations taken in better seeing conditions to detect it. Offset objects in which the peak and majority of the line emission is centred off the BCG account for 4 of the 73 objects. Two of these objects were studied in detail in Hamer et al. (2012) and it was found that the offset line emission was co-spatial with the peak of X-ray emission from the ICM. Another $\sim 15\%$ (10/73) of the sample showed plumes of $\text{H}\alpha$ emission which had a bright peak centred on the BCG but

also showed a clear extent in one preferential direction. For one object (Abell 3880) this plume appears to lead to another object within the VIMOS field of view. However, we note that this object shows spectral absorption lines at $z=0$ clearly identifying it as a foreground star. One possible explanation for plumes is that they are caused by a cooling wake similar to that proposed by Fabian et al. (2001) as the origin of the long filaments in Abell 1795. While in Abell 1795 the resulting $H\alpha$ emission appears as long filaments that we do not see in our plumes, we note that the thickness of the filaments in Abell 1795 is significantly smaller to our mean seeing and all our plumed objects are of comparable or greater redshift than Abell 1795. As such the seeing may be blending the filaments together in our observations producing the plumes that we see. A dedicated study of these objects combining high resolution X-ray and $H\alpha$ observations and comparing to Abell 1795 would be needed to confirm this.

Disturbed objects have line emission that peaks close to the centre of the BCGs continuum peak but also shows extended non-uniform structures at lower surface brightness. There are 13 such objects ($\sim 18\%$) present in the sample which suggests that they are not common amongst the BCG population. It is possible that their prevalence in previous studies is a result of choosing to observe the most extreme systems. Indeed the most common type of morphology seen in our sample are quiescent, simple elliptical systems which are concentrated towards the centre of the BCG. More than 60% of the sample (45/73) show a Quiescent morphology suggesting that major events (such as interactions) which can disturb the gas are relatively rare and that many systems remain undisturbed for long periods of time.

One caveat that must be considered however, is that the observations of the sample are not able to resolve the filamentary structures seen in narrow band imaging of objects such as NGC 5044 (Sun 2012) and NGC 1275 (Conselice et al. 2001). Such filaments are thought to be created when gas is uplifted through AGN feedback as the expanding radio lobe entrains the gas. The filaments in NGC 5044 appear as an extended, non-uniform low surface brightness emission region in our IFU observations and make the system appear disturbed. It is possible that several of the other disturbed systems in our sample are similar to NGC 5044 with filamentary structures throughout their extended emission. We do identify the presence of filaments of $H\alpha$ emission in six objects, Abell 3378, Abell 3998, Abell 85, RXCJ 0120.9-1351, RXCJ 1304.2-3030 and S 805. While it is possible these are related to the filaments seen in narrow-band observations we note that the apparent thickness of these structures is substantially higher than that of those observed with narrow-band filters. While this may be an effect of the poorer seeing of our observations these systems would have solitary long filaments, more like those of Abell 1795 (McDonald & Veilleux 2009) than NGC 1275 or NGC 5044.

The kinematics of the ionised gas is also an important indicator of how disturbed the system is. The mean velocity maps produced by the fitting routine allow us to directly examine the line of sight velocity structure of each object. The velocity structure of the ionised gas in two thirds of the sample appears highly organised within the central regions with most objects showing a continuous velocity gradient across

the brightest region of the $H\alpha$ emission which is typically coincident with the BCG.

To test the significance of this ordered velocity structure we produced a simple model velocity field for each object following a single velocity gradient through the peak of the $H\alpha$ emission and adjusted the PA and gradient to minimise the residuals with the $H\alpha$ velocity map. We define the disorder in the velocity map to be the standard deviation of these residuals ($\sigma_{vel-resid}$). We calculate the strength of the velocity field as the peak-to-peak velocity divided by $\sigma_{vel-resid}$ and the region of the line emission associated with the velocity gradient as having residuals less than $\sigma_{vel-resid}/3$. We compare the strength of the velocity field with the fraction of the $H\alpha$ flux associated to it in Figure 25. More than half our objects (40 out of 73) fall in the upper right hand region of this plot (velocity field strength greater than 3 and more than half the $H\alpha$ flux associated with the velocity gradient) suggesting that the majority of their $H\alpha$ emission is associated with a strong, ordered velocity field. We note that taking a velocity field strength of 2 as the dividing limit places roughly 2/3rds of our objects in this category consistent with our visual inspection. This may not be an unrealistic assumption given that we use a very simple model for the velocity field and it is likely a more sophisticated model would have lower residuals and thus the value of the peak-to-peak velocity divided by the standard deviation would be higher.

The more extended low surface brightness emission present in some systems does not seem to share this ordered line of sight velocity structure, showing a much more random and disturbed distribution of velocities. We would expect to see increased line widths in regions of the gas which is turbulent and disturbed. However, the outer regions in most systems appear to have low line widths which are relatively uniform suggesting that this is not the case. It is important to note that in systems with a large line of sight velocity shift, the coarse spatial sampling can result in artificially broadened lines but we determine this effect will be negligible given the line widths and velocity gradients measured (see Figure 13 and the related discussion).

A number of systems (9/73) showed multiple velocity components in their line emission. Typically these additional components are very broad line emission (8/73) in the very central regions of the BCG on scales comparable to the seeing and are most likely emission from the nuclear regions where gas is feeding onto the AGN. Indeed this is expected as the presence of strong radio sources within the BCGs of many clusters suggest that cold gas is actively being accreted in many systems. More interestingly however, two of our 73 objects showed multiple narrow line components similar to those found in Abell 1664 (Wilman et al. 2006, 2009). These objects were Abell 3574 and NGC 5044 and the two are very different. NGC 5044 shows a double component in its core regions (in addition to a broad line) while Abell 3574 has a double component ring of gas which surrounds its core. The full interpretation of these double component systems requires more data but suggests the ionised gas may be much more disturbed in a small fraction ($\sim 3\%$) of systems.

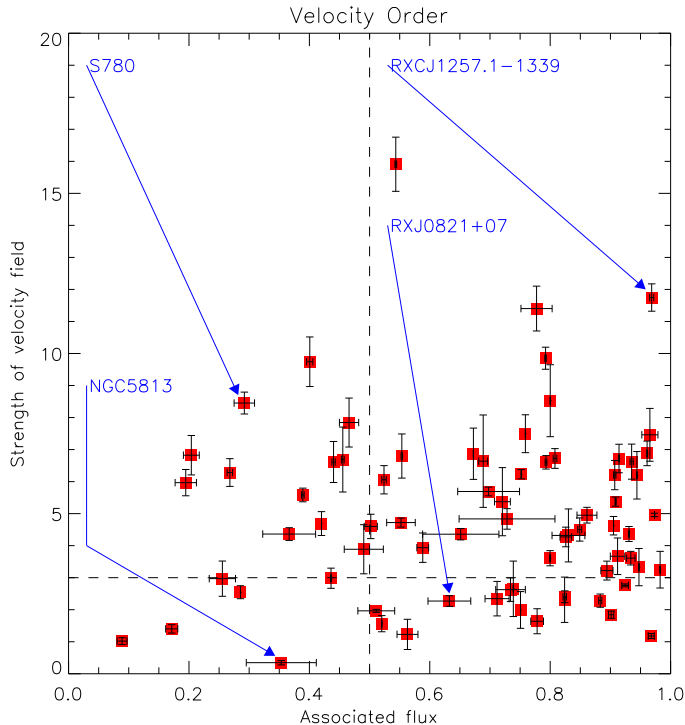


Figure 25. A comparison of the strength of the ordered structure in the velocity field (measured in standard deviations of the residuals) against the flux associated to the velocity field as a fraction of the total flux. The horizontal dashed line separates the strength of the velocity fields at $3 \sigma_{vel-resid}$ while the vertical dashed line separates the associated flux at 50%. One object from each quadrant is labelled to allow comparison with the velocity maps. More than half our objects fall into the upper right hand quadrant suggesting that the majority of their H α emission is associated with a strong velocity field.

6.2 What role does the BCG play in the cooling of gas from the ICM?

In Hamer et al. (2012) we studied several objects in detail and determined that (for that subsample at least) the offset H α emission was coincident with the peak of the X-ray emission from the ICM. Our analysis of the X-ray – optical line offsets in Section 4.1.1 would appear to confirm this result as at most 3 of the 24 objects studied showed a potential significant offset of the H α emission from the peak of the X-rays. The analysis of the structure of the full sample presented in this paper does identify several more objects which show an offset between the H α and continuum peaks. Hamer et al. (2012) were able to associate the offset peak with evidence of gas sloshing (Johnson et al. 2010, 2012; ZuHone et al. 2010, 2011; Blanton et al. 2011) in the ICM as well as continued cooling at the position offset from the BCG. However, in these additional systems the multi-wavelength data necessary to make this determination (deep X-ray imaging and/or resolved CO detections) were not available so we cannot say if this is true for all offset systems. Despite this it is apparent that offset line emission, while still rare (4/73 systems), is more common than initially thought and may thus play a more significant role in the evolution of the cluster core.

In a small number of objects the stellar continuum was sufficiently bright to allow us to extract the stellar kinematics from the BCG. Comparing the maps of the stellar

kinematics to the gas kinematics it is apparent that the line of sight velocity structure of the stellar component is negligible, in stark contrast to the high velocity, ordered structures seen in the gas component. This suggests that the kinematics of the line emitting gas are decoupled from those of the stellar population of the BCG implying that cooling gas is not necessarily related to the BCG even when cooling is occurring at the BCGs location.

This raises the question, if the gas is relaxed and stable why is it not forming stars and producing a stellar velocity field that matches that seen in the gas? Hydra-A has one of the most extreme H α velocity fields in our sample, the gas is contained in a rotating disc that is known to be forming stars at a rate of $2-3 \text{ M}_{\odot} \text{ yr}^{-1}$ (Donahue et al. 2011; Hoffer et al. 2012). If the disc is responsible for feeding the AGN (Hamer et al. 2014) then it must have been stable for the lifetime of the outburst ($\sim 10^8 \text{ yr}$ for an accretion rate of $0.1-0.25 \text{ M}_{\odot} \text{ yr}^{-1}$ and a total accreted mass of 10^7 M_{\odot} Wise et al. 2007). Assuming the star formation rate is constant during this time then the total mass of stars the disc will form is $3 \times 10^8 \text{ M}_{\odot}$, which is a negligible fraction of the typical stellar mass of a BCG ($10^{11}-10^{12} \text{ M}_{\odot}$). Thus the stellar velocity field we observe is dominated by the bulk of the stellar component formed in previous generations of star formation, which need not share the current velocity field of the gas.

6.3 What role does the cold gas play in fuelling feedback?

Mechanical feedback from AGN outbursts within BCGs is commonly believed to be the main contributing factor preventing the catastrophic cooling of gas in cluster cores. A requirement for this process to work effectively is a mechanism of self-regulation to maintain the balance between cooling gas and feedback. Thus a link must be established between the gas cooling (on kpc scales) and its supply to the super massive black hole which results in the activation of the AGN once cooling reaches a critical threshold. In Hamer et al. (2014) we show that one system (Hydra-A) in our sample has kinematics in the ionised gas that clearly trace the presence of an extended ($\sim 5 \text{ kpc}$) disc of cold, molecular and atomic gas in the core of the system. Such a disc can potentially channel gas to the centre of the system relatively quickly allowing it to fuel the AGN. The confirmation of the disc in Hydra-A was made possible thanks to the system being relatively close by, having a large velocity shift and being viewed at close to edge on. Our analysis of the velocity gradients suggest that this sample is consistent with a sample of rotating discs similar to the one seen in Hydra-A and Figure 25 suggests that $> 50\%$ of systems have the bulk of their H α emission associated with this velocity gradient. However, confirming the presence of a disc within each individual system requires a more sophisticated, case-by-case analysis (which will be presented in paper II, Hamer *et al.* in prep.).

From the velocity dispersion maps it is apparent that the linewidths are typically broadest towards the peak of the H α emission which is consistent with this position having the longest line of sight through the gas and thus sampling the widest distribution of velocity components. We note that

the FWHM in the central regions is typically on the order of the velocity shift across the whole system suggesting it originates from the extended gas in the ionised nebula. However, 8 of the 73 systems show some evidence of very broad (FWHM on the order of 2355 km s^{-1} or more) components localised (on scales of the seeing) to their central regions. This can be interpreted as emission from an AGN suggesting that the AGN in these systems are actively accreting. However, we note that the broad component is typically very weak and the scales of the seeing are much greater than those expected for an AGN making this difficult to confirm without deeper, higher resolution observations.

6.4 The source of the ionisation

The maps of the $[\text{NII}]_{\lambda 6583}/\text{H}\alpha$ ratio produced in Section 4 show a wide variation in value between objects. More interestingly perhaps the ratio is seen to vary within a given object suggesting that the source of the ionising radiation may not be uniformly distributed. Likewise it shows no trend radiating from a single point suggesting that the dominant source of the ionisation is not a single localised source but is more homogeneously spread throughout the cluster core. Given the extended nature of these sources, and the variation seen within different regions of individual objects, it seems unlikely that an AGN positioned at the centre of the BCG could be responsible for the ionisation throughout the nebula. However, the trend that more compact objects typically show higher line ratios than extended objects does suggest that AGN still play a role in the excitation of the gas even if it is not dominant. We therefore conclude that other, more diffuse, sources must also be significantly contributing to the excitation of the gas.

Star formation is (along with AGN) the most commonly invoked excitation mechanism and has the potential to be distributed throughout the BCG. For the few objects (just 7, less than 10% of our sample) in which we had access to the $\text{H}\beta$ and $[\text{OIII}]$ lines and we perform a BPT analysis (Baldwin et al. 1981) to look for evidence of excitation by star formation. This analysis suggests that the gas in all of these objects is predominantly ionised by non stellar processes and is, as such, inconsistent with excitation by star formation. Analysis with other spectral indicators (the $[\text{NII}/\text{H}\alpha]$, $[\text{SII}]/\text{H}\alpha$ and $[\text{OI}]/\text{H}\alpha$ emission line ratios) also indicates that non stellar processes dominate the ionisation of the gas in most systems. These line ratio diagnostics are all consistent and rule out star formation as the dominant excitation mechanism in the vast majority of objects in our sample.

Shocks also seem to be an appropriate candidate for a distributed excitation mechanism however, we note that the velocity dispersion maps do not show evidence of shocks of sufficient velocity in the extended regions (seen as enhanced line widths) to produce the line ratios seen. Comparison of the measured line ratios with those predicted by shock models also shows no consistent evidence that shocks are a major contributor to the gas excitation for the objects in our sample. The most likely candidate then is particle heating caused by a reconnection of the gas phases (Ferland et al. 2009; Fabian et al. 2011). This can also explain the presence of ionised emission in offset objects which have bright line emission coincident with their X-ray peaks (Hamer et al.

2012) but are displaced from the bulk of the BCGs stellar component. We note however, that AGN still play a role in the excitation of the gas and indeed other sources of excitation, such as cosmic rays, shocks and star formation, are likely also contributing to the overall ionisation state of the gas. All of these processes, and their interaction, need to be considered to fully understand the excitation of the gas.

7 CONCLUSIONS

We have analysed the morphology, kinematics and excitation of the optical line emitting nebula in a sample of 73 BCGs with the VIMOS IFU. Our goal has been to develop a more complete understanding of the nature of these ionised nebula, what role they play in the physical processes affecting the cluster core and what they can tell us about the interaction of the BCG with the ICM. Throughout this paper we present a number of results that can be used to address these issues which we will now summarise.

- 1) We classify our objects into five categories based on the morphology of their optical line emitting nebula. While we do find that some objects appear disturbed (13/73), the majority of our sample have elliptical, centrally concentrated morphologies (45/73) suggesting that the gas is in a relatively quiescent, undisturbed state (Section 4.1).
- 2) We detect a number of systems in which the peak of the line emission is offset from the BCG. Some of these objects where already known and they are typically found in systems in which the ICM is “sloshing” (Hamer et al. 2012) suggesting that the gas in the ICM has been disturbed. The number of these objects in our sample is small (4/73) suggesting that such major disturbances are rare (Section 4.1.1).
- 3) The extent of the nebulae varies quite considerably between objects but can extend out over several tens of kpc in some objects. While previous narrow band studies suggest we might expect to see complex filamentary structures in these extended regions we are unable to resolve the filamentary structures in most objects. However, in some systems we do detect isolated filaments which may be related and only detectable because they are isolated and thus not blurred with other filaments as a result of the seeing (Section 4.1).
- 4) The majority of our objects show gas with a high line of sight velocity in the form of a velocity shift that runs across the ionised nebula. This velocity structure is visible in the velocity maps (Section 4.2) but it is also visible in the channel maps (Section 3) where the position of the gas can be seen to vary smoothly with channel.
- 5) The velocity structure of the gas in most systems appears quite ordered and typically varies little away from the main velocity shift that runs across the nebula (Section 4.2). This is especially apparent when we consider the strength of this velocity shift against the fraction of the $\text{H}\alpha$ flux that it represents (Figure 25) and we find that in 40/73 of our objects more than half the $\text{H}\alpha$ flux is associated with a strong velocity shift.
- 6) When measuring the gradient of the velocity shifts observed we find that the gradient shows a clear relationship with the separation of the maximum and minimum

velocities but none with the magnitude of the velocity difference. The relationship we measure is inconsistent with random motions suggesting that the measured velocity along the velocity shift is a function of distance from the BCG (Section 4.2).

- 7) The velocity dispersion maps of the gas show that the linewidth typically peaks towards the centre of the nebula but is otherwise low ($\sim 100\text{--}200\text{ km s}^{-1}$) in the more extended regions (Section 4.2.1).
- 8) We show that the total $\text{H}\alpha$ luminosity and extent of the line emitting nebula are well correlated. This coupled with the fact the velocity of the gas appears related to the distance from the core of the BCG suggests that the magnitude of the velocity shift should be related to the $\text{H}\alpha$ luminosity. By assuming that differences in the velocity shift at a given luminosity are caused by inclination we are able to reproduce an inclination distribution consistent with what is expected for a random sample such as this (Section 4.3).
- 9) By measuring the stellar kinematics from the NaD absorption feature we show that the stars do not have a strong velocity shift, or any obvious velocity structure, like those seen in the gas. This suggests that the velocity field of the stars is dominated by random motions and is thus decoupled from the gas component (Section 4.4).
- 10) The maps of measured line ratio ($[\text{NII}]/\text{H}\alpha$) show a considerable amount of variation within each nebula. However, importantly they show no obvious trend or ordered structure suggesting that the source of the excitation is not localised to a specific region of the nebula (Section 5.3).
- 11) By comparing the global $[\text{NII}]/\text{H}\alpha$ ratio for each object to the extent of the line emitting nebula we find a loose trend that objects with a smaller extent have higher line ratios. This implies that more compact objects are preferentially excited by more energetic sources suggesting that AGN may be playing some role in the excitation of the gas, at least in the central regions of the BCG (Section 5.3).
- 12) We perform line ratio diagnostics on the gas to determine the source of the excitation and find that for most objects in our sample it is inconsistent with star formation. We are also unable to find any consistent evidence that shocks are playing a significant role in the excitation (Section 5.4).

These results allow us to draw several conclusions about the nature of the ionised gas within cluster cores. The fact that the majority of objects have a $\text{H}\alpha$ morphology which suggests the gas is relatively quiescent (Point 1) suggests that the gas is typically not highly disturbed within most systems. Similarly while many objects do show gas at high velocities relative to the BCG the structures of the high velocity components appear well ordered varying smoothly from blue shifted to red shifted across the nebula (Points 4 and 5). If the gas were disturbed then we might expect to see evidence of shocks, caused by moving gas colliding with gas moving at different velocities, in the velocity dispersion maps. However, the velocity dispersion maps do not show evidence of enhanced line widths away from the central regions (Point 7) suggesting that random motions are not significant within the extended nebula. We do see a number of systems

with disturbed morphologies but these objects account for less than 20% (13/73) of our sample (Point 1). These objects also typically show a more disordered velocity structure in the extended, low surface brightness gas (e.g. Abell 3581, NGC 5044, RXJ 0338+09). We also identify several objects with offset emission similar to those studied in Hamer et al. (2012) (Point 2). While the nature of these objects clearly suggests they are disturbed they account for a very small fraction of the overall population (just 4/73 of our objects). While a few of these disturbed systems did show evidence of a possible interaction with another cluster member (e.g. RXJ 0338+09) it was not a common factor in all disturbed systems. However, we cannot rule out the possibility that the gas was disturbed by a previous fly-by interaction with a galaxy that has since left the VIMOS field of view. Despite these systems the majority of the BCGs in our sample have line emitting nebula which are quiescent and ordered in both their morphology and kinematics suggesting that the gas is relaxed rather than disturbed within the bulk of the population.

The identification of several more objects with offset emission similar to those studied in Hamer et al. (2012) (Point 2) suggests that such offsets may be more common than previously thought. It is possible that the multiple velocity components seen in the cores of objects like Abell 1664 and NGC 5044 may be a related phenomena but a more detailed analysis of these objects is needed to confirm such a link. If these additional offset systems are similar to those presented in Hamer et al. (2012) then it suggests that sloshing may be present (albeit at different levels) in all systems and strengthens the conclusion that the BCG is not required for cooling gas to condense from the ICM despite the strong spatial correlation with the cluster core (Peres et al. 1998; Sanderson et al. 2009; Hudson et al. 2010). The fact that the gas does not share the kinematics of the stars (Point 9) also suggests that the stellar component of the BCG and the gas that make up the nebula are not linked and that the gas need not have cooled at its current location within the BCG. The large extent of some of the objects in our sample and the filaments seen may also support this (Point 3). While these structures may be the result of gas which has been removed from the BCG (either entrained by radio jets or blown out by winds) the fact that we see gas cooling away from the BCG in the offset systems suggests a different possibility. These filaments and extended regions could be the result of gas which has cooled from the ICM away from the BCG and is now falling towards it. Our current data do not allow us to test this hypothesis but with new instrumentation (MUSE, ALMA etc.) it should be possible.

The shift from positive to negative velocities across the brightest regions of the $\text{H}\alpha$ seen in the majority of the systems in our sample (Point 5) is similar to the velocity structure found in Hydra-A (Hamer et al. 2014). We also note that the velocity gradients show evidence of being dependant only on the distance over which they are measured (Point 6). Such a velocity structure could be the result of a disc like rotation similar to that seen in Hydra-A and our estimations of the inclinations from the magnitude of the velocity shift at a given luminosity (Point 8) are consistent with the expected inclinations of a random sample of discs. This suggests that kpc scale discs may be common in cluster cores however, we note that other phenomena can also

produce velocity structures similar to those seen (bi-polar outflows for example) so a more detailed study is needed to confirm the presence of discs. If discs are common this would provide a potential link between the gas cooling on kpc scales and the fuelling of the central black hole. The fact that the stellar and gas kinematics are decoupled (Point 9) seems to support this interpretation. If the gas is stable as it appears to be, then it should form stars which over the lifetime of the cluster would produce a stellar velocity field matching that of the gas. Thus gas discs such as the one seen in Hydra-A must be a transient phenomena and the gas must be disturbed or consumed before it can form a sufficient number of stars to affect the overall stellar velocity field. The main questions remaining then are whether the velocity structure seen in the gas within most BCGs relates to a disc and if it does, is the disc jet alignment seen in Hydra-A (Hamer et al. 2014) unique to that system or common to all BCGs with gas discs. We will address these questions in the next paper in this series by presenting a detailed analysis of the kinematics of each individual object and comparing to tracers of the AGN outbursts at radio and X-ray wavelengths.

From our study of the gas excitation we conclude that a distributed process is the main component of the gas ionisation as the line ratio maps show significant variation within a given object but no obvious trend from a given point to suggest a localised source (Point 10). Despite this the global line ratios for each object are higher in more compact objects (Point 11) which suggests that an energetic source at the centre of the emission (most likely an AGN) is contributing to the excitation of the gas but is not the dominant component in the most extended regions. Through the use of line ratio diagnostics we can rule out star formation as the dominant distributed excitation source (Point 12). The line ratio diagnostics also show no consistent evidence for shocks (Point 12) and the narrow linewidths seen in the extended emission suggest that the fast shocks needed to excite the gas to the levels seen are not present (Point 7). As such we must appeal to particle heating (Ferland et al. 2009; Fabian et al. 2011) to explain the excitation in the extended nebula, though AGN (and likely other processes) still play a role.

This work represents a significant step forward in the study of the interplay between the Brightest Cluster Galaxy, the cold gas in the cluster core and the cooling of the hot gas in the ICM. We have shown that despite the high velocity of the ionised gas it typically forms simple and ordered spatial and kinematics structures. Such simple structures clearly favour a situation in which the gas has cooled from the surrounding medium and has remained undisturbed by interactions with other cluster members for some time.

ACKNOWLEDGEMENTS

SLH acknowledges support from the European Research Council for Advanced Grant Program num 339659–MUSICOS. FC acknowledges the support from the European Research Council for Advanced Grant Program num 267399–Momentum. AMS acknowledges an STFC Advanced Fellowship. ACE acknowledges support from STFC grant ST/I001573/1.

Based on observations made with ESO Telescopes at the La Silla or Paranal Observatories under programme ID 080.A-0224 and 082.B-0671

REFERENCES

- Abell G. O., 1958, *ApJS*, 3, 211
 Allen M. G., Groves B. A., Dopita M. A., Sutherland R. S., Kewley L. J., 2008, *ApJS*, 178, 20
 Bahcall N. A., 1999, in Dekel A., Ostriker J. P., eds, *Formation of Structure in the Universe Clusters and super-clusters of galaxies*. p. 135
 Baldwin J. A., Phillips M. M., Terlevich R., 1981, *PASP*, 93, 5
 Bica E., Pastoriza M. G., da Silva L. A. L., Dottori H., Maia M., 1991, *AJ*, 102, 1702
 Blanton E. L., Randall S. W., Clarke T. E., Sarazin C. L., McNamara B. R., Douglass E. M., McDonald M., 2011, *ApJ*, 737, 99
 Böhringer H., Schuecker P., Guzzo L., Collins C. A., Voges W., Cruddace R. G., Ortiz-Gil A., Chincarini G., De Grandi S., Edge A. C., MacGillivray H. T., Neumann D. M., Schindler S., Shaver P., 2004, *A&A*, 425, 367
 Böhringer H., Werner N., 2009, *ArXiv:0907.4277*
 Bower R. G., Benson A. J., Malbon R., Helly J. C., Frenk C. S., Baugh C. M., Cole S., Lacey C. G., 2006, *MNRAS*, 370, 645
 Braine J., Dupraz C., 1994, *A&A*, 283, 407
 Burns J. O., Rhee G., Owen F. N., Pinkney J., 1994, *ApJ*, 423, 94
 Cappellari M., Emsellem E., 2004, *PASP*, 116, 138
 Cavagnolo K. W., Donahue M., Voit G. M., Sun M., 2008, *ApJ*, 683, L107
 Cavagnolo K. W., Donahue M., Voit G. M., Sun M., 2009, *ApJS*, 182, 12
 Churazov E., Sunyaev R., Forman W., Böhringer H., 2002, *MNRAS*, 332, 729
 Conselice C. J., Gallagher III J. S., Wyse R. F. G., 2001, *AJ*, 122, 2281
 Courteau S., 1997, *AJ*, 114, 2402
 Crawford C. S., Allen S. W., Ebeling H., Edge A. C., Fabian A. C., 1999, *MNRAS*, 306, 857
 Crawford C. S., Fabian A. C., 1992, *MNRAS*, 259, 265
 Crawford C. S., Gandhi P., Fabian A. C., Wilman R. J., Johnstone R. M., Barger A. J., Cowie L. L., 2002, *MNRAS*, 333, 809
 Croton D. J., Springel V., White S. D. M., De Lucia G., Frenk C. S., Gao L., Jenkins A., Kauffmann G., Navarro J. F., Yoshida N., 2006, *MNRAS*, 365, 11
 David L. P., Lim J., Forman W., Vrtillek J., Combes F., Salome P., Edge A., et al. 2014, *ApJ*, 792, 94
 David L. P., O’Sullivan E., Jones C., Giacintucci S., Vrtillek J., Raychaudhury S., Nulsen P. E. J., Forman W., Sun M., Donahue M., 2011, *ApJ*, 728, 162
 Denicoló G., Terlevich R., Terlevich E., 2002, *MNRAS*, 330, 69
 Donahue M., de Messières G. E., O’Connell R. W., Voit G. M., Hoffer A., McNamara B. R., Nulsen P. E. J., 2011, *ApJ*, 732, 40
 Dopita M. A., Kewley L. J., Heisler C. A., Sutherland R. S., 2000, *ApJ*, 542, 224

- Ebeling H., Edge A. C., Allen S. W., Crawford C. S., Fabian A. C., Huchra J. P., 2000, *MNRAS*, 318, 333
- Ebeling H., Edge A. C., Bohringer H., Allen S. W., Crawford C. S., Fabian A. C., Voges W., Huchra J. P., 1998, *MNRAS*, 301, 881
- Ebeling H., Voges W., Bohringer H., Edge A. C., Huchra J. P., Briel U. G., 1996, *MNRAS*, 281, 799
- Edge A. C., 2001, *MNRAS*, 328, 762
- Edwards L. O. V., Robert C., Mollá M., McGee S. L., 2009, *MNRAS*, 396, 1953
- Fabian A. C., 1994, *ARA&A*, 32, 277
- Fabian A. C., Hu E. M., Cowie L. L., Grindlay J., 1981, *ApJ*, 248, 47
- Fabian A. C., Johnstone R. M., Sanders J. S., Conselice C. J., Crawford C. S., Gallagher III J. S., Zweibel E., 2008, *Nature*, 454, 968
- Fabian A. C., Sanders J. S., Ettori S., Taylor G. B., Allen S. W., Crawford C. S., Iwasawa K., Johnstone R. M., 2001, *MNRAS*, 321, L33
- Fabian A. C., Sanders J. S., Williams R. J. R., Lazarian A., Ferland G. J., Johnstone R. M., 2011, *MNRAS*, 417, 172
- Ferland G. J., Fabian A. C., Hatch N. A., Johnstone R. M., Porter R. L., van Hoof P. A. M., Williams R. J. R., 2009, *MNRAS*, 392, 1475
- Hamer S. L., Edge A. C., Swinbank A. M., Oonk J. B. R., Mittal R., McNamara B. R., Russell H. R., Bremer M. N., Combes F., Fabian A. C., Nesvadba N. P. H., O'Dea C. P., Baum S. A., Salomé P., Tremblay G., Donahue M., Ferland G. J., Sarazin C. L., 2014, *MNRAS*, 437, 862
- Hamer S. L., Edge A. C., Swinbank A. M., Wilman R. J., Russell H. R., Fabian A. C., Sanders J. S., Salomé P., 2012, *MNRAS*, p. 2520
- Hatch N. A., Crawford C. S., Fabian A. C., Johnstone R. M., 2005, *MNRAS*, 358, 765
- Hatch N. A., Crawford C. S., Johnstone R. M., Fabian A. C., 2006, *MNRAS*, 367, 433
- Heckman T. M., 1980, *A&A*, 87, 142
- Heckman T. M., Baum S. A., van Breugel W. J. M., McCarthy P., 1989, *ApJ*, 338, 48
- Heckman T. M., Lehnert M. D., Strickland D. K., Armus L., 2000, *ApJS*, 129, 493
- Hlavacek-Larrondo J., Fabian A. C., Edge A. C., Ebeling H., Allen S. W., Sanders J. S., Taylor G. B., 2013, *MNRAS*, 431, 1638
- Hoffer A. S., Donahue M., Hicks A., Barthelmy R. S., 2012, *ApJS*, 199, 23
- Hogan M. T., Edge A. C., Hlavacek-Larrondo J., Grainge K. J. B., Hamer S. L., Mahony E. K., Russell H. R., Fabian A. C., McNamara B. R., Wilman R. J., 2015, *MNRAS*, 453, 1201
- Hudson D. S., Mittal R., Reiprich T. H., Nulsen P. E. J., Andernach H., Sarazin C. L., 2010, *A&A*, 513, A37+
- Jacoby G. H., Hunter D. A., Christian C. A., 1984, *ApJS*, 56, 257
- Jiménez-Vicente J., Castillo-Morales A., Mediavilla E., Battaner E., 2007, *MNRAS*, 382, L16
- Johnson R. E., Markevitch M., Wegner G. A., Jones C., Forman W. R., 2010, *ApJ*, 710, 1776
- Johnson R. E., Zuhone J., Jones C., Forman W. R., Markevitch M., 2012, *ApJ*, 751, 95
- Johnstone R. M., Fabian A. C., 1988, *MNRAS*, 233, 581
- Johnstone R. M., Fabian A. C., Nulsen P. E. J., 1987, *MNRAS*, 224, 75
- Katz N., White S. D. M., 1993, *ApJ*, 412, 455
- Kewley L. J., Dopita M. A., 2002, *ApJS*, 142, 35
- Kewley L. J., Dopita M. A., Sutherland R. S., Heisler C. A., Trevena J., 2001, *ApJ*, 556, 121
- Kewley L. J., Groves B., Kauffmann G., Heckman T., 2006, *MNRAS*, 372, 961
- Loewenstein M., 2004, *Origin and Evolution of the Elements*, p. 422
- Markevitch M., Vikhlinin A., 2007, *Phys. Rep.*, 443, 1
- McDonald M., Veilleux S., 2009, *ApJ*, 703, L172
- McDonald M., Veilleux S., Mushotzky R., 2011, *ApJ*, 731, 33
- McNamara B. R., Jaffe W., 1994, *A&A*, 281, 673
- McNamara B. R., Nulsen P. E. J., 2007, *ARA&A*, 45, 117
- McNamara B. R., Nulsen P. E. J., Wise M. W., Rafferty D. A., Carilli C., Sarazin C. L., Blanton E. L., 2005, *Nature*, 433, 45
- McNamara B. R., O'Connell R. W., 1989, *AJ*, 98, 2018
- More J., Wright S., 1993, *Optimization Software Guide*. *Frontiers in Applied Mathematics*, Society for Industrial and Applied Mathematics
- O'Dea C. P., Baum S. A., Maloney P. R., Tacconi L. J., Sparks W. B., 1994, *ApJ*, 422, 467
- O'Dea K. P., Quillen A. C., O'Dea C. P., Tremblay G. R., Snios B. T., Baum S. A., Christiansen K., Noel-Storr J., Edge A. C., Donahue M., Voit G. M., 2010, *ApJ*, 719, 1619
- Osterbrock D. E., 1989, *Astrophysics of gaseous nebulae and active galactic nuclei*
- Owen F. N., Ledlow M. J., 1997, *ApJS*, 108, 41
- Peres C. B., Fabian A. C., Edge A. C., Allen S. W., Johnstone R. M., White D. A., 1998, *MNRAS*, 298, 416
- Peterson J. R., Fabian A. C., 2006, *Phys. Rep.*, 427, 1
- Peterson J. R., Kahn S. M., Paerels F. B. S., Kaastra J. S., Tamura T., Bleeker J. A. M., Ferrigno C., Jernigan J. G., 2003, *ApJ*, 590, 207
- Peterson J. R., Paerels F. B. S., Kaastra J. S., Arnaud M., Reiprich T. H., Fabian A. C., Mushotzky R. F., Jernigan J. G., Sakellou I., 2001, *A&A*, 365, L104
- Pettini M., Pagel B. E. J., 2004, *MNRAS*, 348, L59
- Russell H. R., McNamara B. R., Edge A. C., Hogan M. T., Main R. A., Vantyghem A. N., 2013, *MNRAS*, 432, 530
- Salomé P., Combes F., 2003, *A&A*, 412, 657
- Salomé P., Combes F., Revaz Y., Downes D., Edge A. C., Fabian A. C., 2011, *A&A*, 531, A85+
- Sanders J. S., Fabian A. C., Taylor G. B., 2009, *MNRAS*, 396, 1449
- Sanderson A. J. R., Edge A. C., Smith G. P., 2009, *MNRAS*, 398, 1698
- Sanderson A. J. R., O'Sullivan E., Ponman T. J., 2009, *MNRAS*, 395, 764
- Sarazin C. L., O'Connell R. W., McNamara B. R., 1992, *ApJ*, 397, L31
- Simkin S. M., 1979, *ApJ*, 234, 56
- Singh R., van de Ven G., Jahnke K., Lyubenova M., Falcón-Barroso J., Alves J., Cid Fernandes R., Galbany L., et al. 2013, *A&A*, 558, A43
- Sun M., 2012, *New Journal of Physics*, 14, 045004
- Tamura T., Kaastra J. S., Peterson J. R., Paerels F. B. S., Mittaz J. P. D., Trudolyubov S. P., Stewart G., Fabian

- A. C., Mushotzky R. F., Lumb D. H., Ikebe Y., 2001, A&A, 365, L87
- Tempel E., Stoica R. S., Saar E., 2013, MNRAS, 428, 1827
- Veilleux S., Cecil G., Bland-Hawthorn J., 2005, ARA&A, 43, 769
- Voigt L. M., Fabian A. C., 2004, MNRAS, 347, 1130
- Werner N., Oonk J. B. R., Sun M., Nulsen P. E. J., Allen S. W., Canning R. E. A., Simionescu A., Hoffer A., Connor T., Donahue M., Edge A. C., Fabian A. C., von der Linden A., Reynolds C. S., Ruszkowski M., 2014, MNRAS, 439, 2291
- Wilman R. J., Edge A. C., McGregor P. J., McNamara B. R., 2011, MNRAS, 416, 2060
- Wilman R. J., Edge A. C., Swinbank A. M., 2006, MNRAS, 371, 93
- Wilman R. J., Edge A. C., Swinbank A. M., 2009, MNRAS, 395, 1355
- Wise M. W., McNamara B. R., Nulsen P. E. J., Houck J. C., David L. P., 2007, ApJ, 659, 1153
- ZuHone J. A., Markevitch M., Johnson R. E., 2010, ApJ, 717, 908
- ZuHone J. A., Markevitch M., Lee D., 2011, ApJ, 743, 16